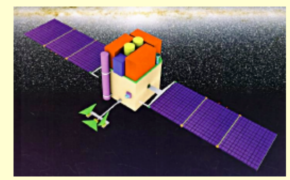
What Thermonuclear X-ray Bursts can tell us about Neutron Stars



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Outline

- Neutron Stars: why do we care?
- Thermonuclear Bursts: why do we care?
- Neutron Stars: Mass, Radius and Spin:
    - a. Continuum Spectroscopy of Bursts
    - b. Spectral Lines from Bursts
    - c. Timing Properties of Bursts
- Neutron Star Atmosphere: Thermonuclear Flame Spreading
- Future Prospects and Conclusions

Neutron Star

Neutron star vs. a city

Figure courtesy M. Coleman Miller

Radius ~ 10 - 20 km

Mass ~ 1.4 - 2.0 solar mass

Core density ~ 5 -10 times the nuclear density

Magnetic field ~ 10^{7} - 10^{15} G

Spin frequency (in some binary stellar systems) ~ 300 - 600 Hz

Some of the most extreme conditions of the universe exist in neutron stars.

Neutron Star: Surface and Interior

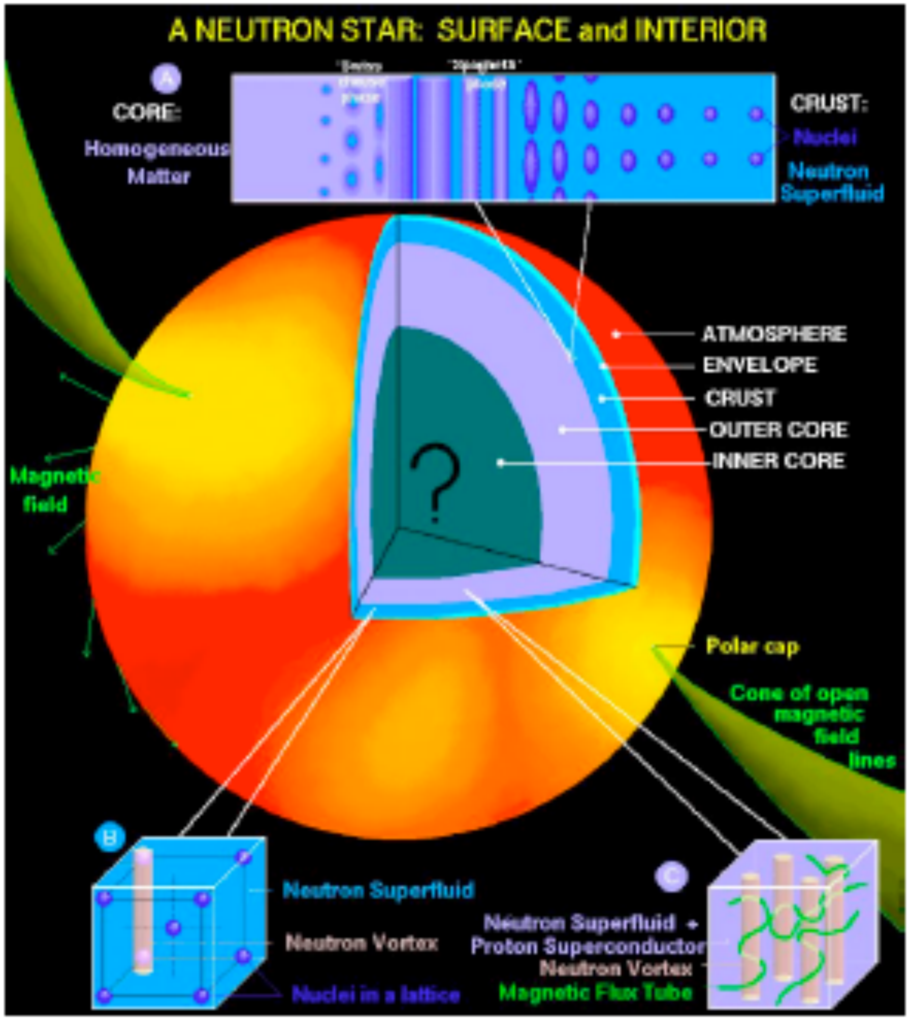


Figure courtesy D. Page.

Core density ≥ nuclear density

⇓

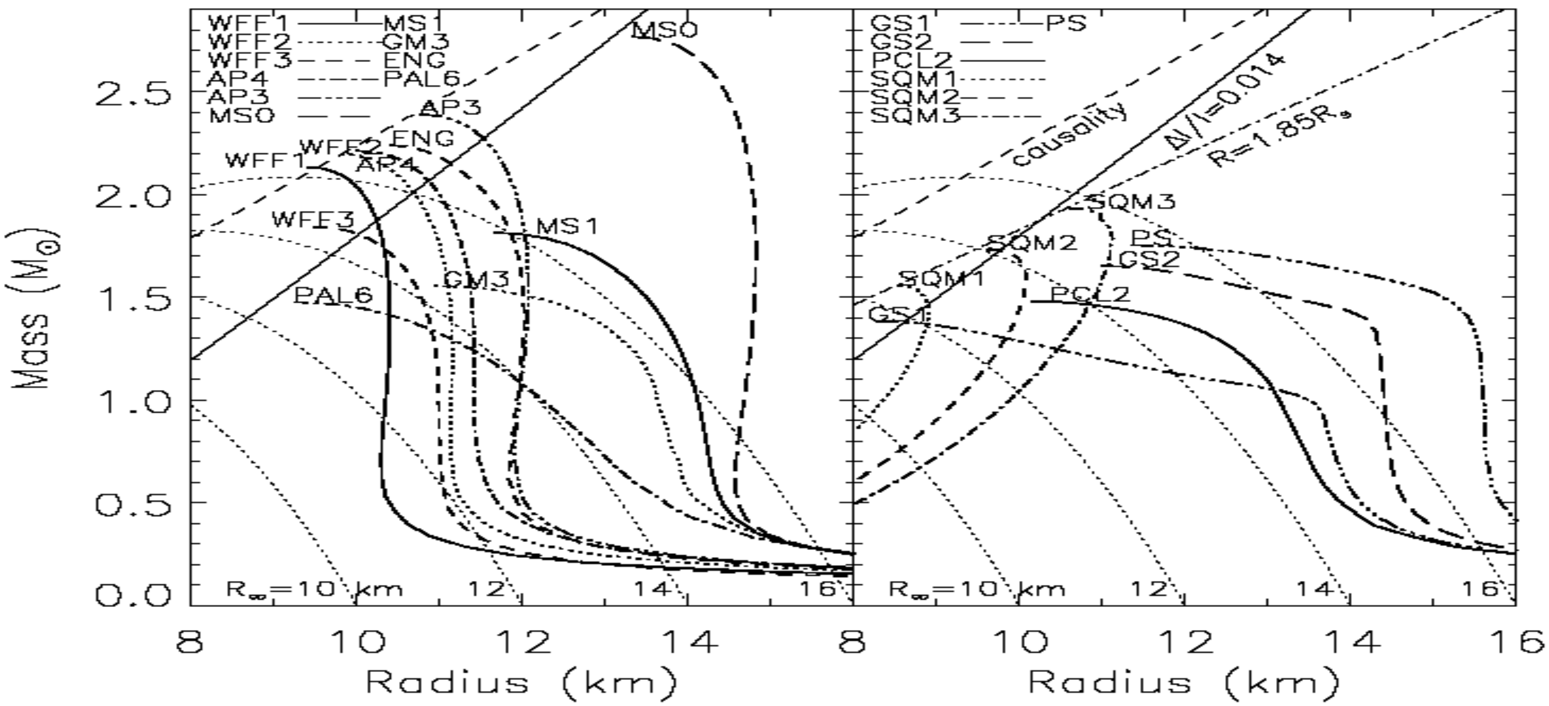
Exotic matter???

No terrestrial experiments seem possible at such high densities and low (comparatively) temperatures.

Many equation of state (EOS) models for the neutron star core matter are available in the literature. We need to constrain these models by observing neutron stars.

The constituents of neutron star interiors remain a mystery after 35 years.

How to constrain EOS models?



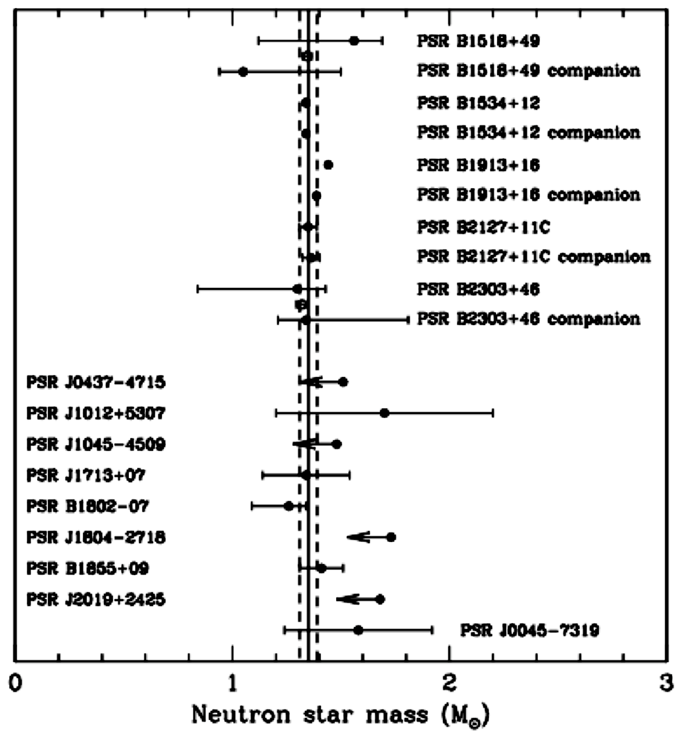
Lattimer & Prakash (2001)

Mass, radius and **spin frequency** of the **same** neutron star are to be measured in order to constrain **equation of state** models.

Neutron Star: Spin, Mass and Radius

Spin frequency measurement: Observation of spin-induced periodic variation of stellar surface intensity.

Mass measurement: Measurements of relativistic orbital effects of binary stellar systems that contain radio pulsars. ⟷
But radius?

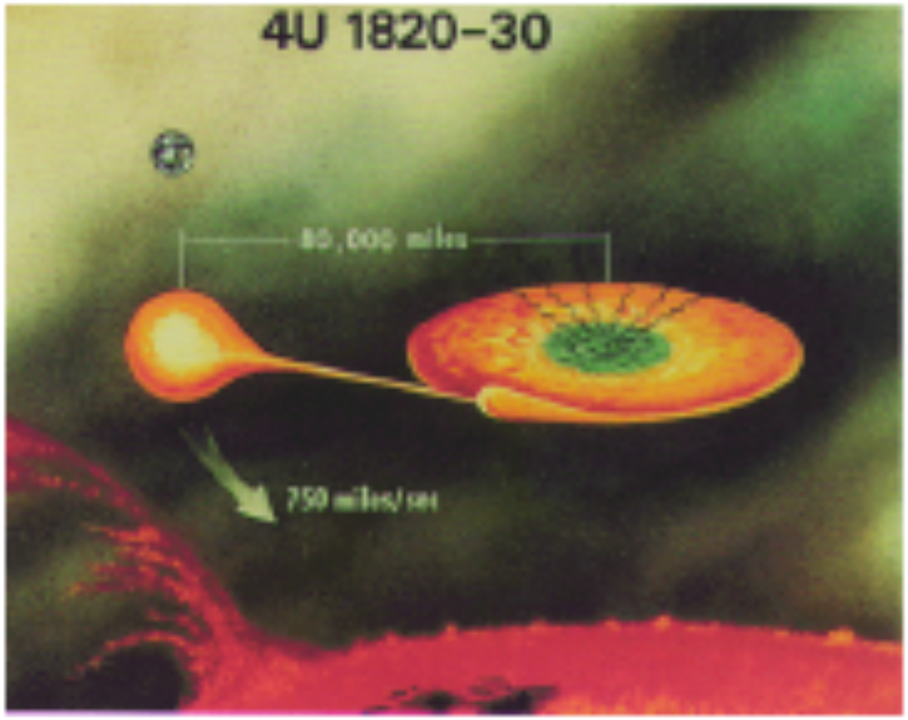


Thorsett and Chakrabarty 1999

Radius measurement: Hardest! To do it with 5% accuracy, we need to measure the size of the neutron star with less than 1 km error, while the star is ~10^{16} km away! Plagued with systematic uncertainties.

We need to study a neutron star surface phenomenon, that will allow us to measure spin, mass and radius of the star.

Low Mass X-ray Binary (LMXB)

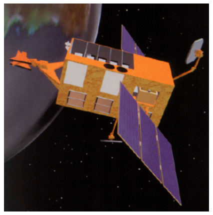


Courtesy: http://heasarc.gsfc.nasa.gov/

Primary star: neutron star, or black hole.
Secondary companion: mass < 1 solar mass; main sequence star, white dwarf, or red giant star.
Accretion: via a disk.
X-rays: from inner accretion disk, and from the neutron star surface.
Age: ~ 10^9 years.
Neutron star magnetic field: 10^7-10^9 G.

X-rays from neutron star surfaces can give important information about these stars.

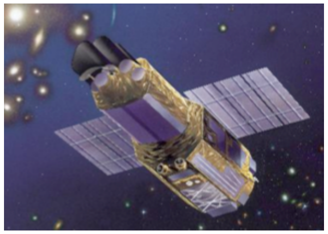
Some Current X-ray Missions



Rossi X-ray Timing Explorer (RXTE): For PCA: Energy range ~ 2 - 60 keV, Effective area (@ 6 keV) ~ 5000 cm^2, Energy resolution (@ 2.5 keV) ~ 725 eV, Angular resolution ~ 1^o, Time resolution ~ 1 microsecond.

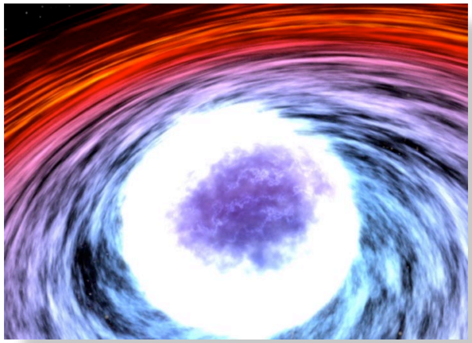
Chandra: Energy range ~ 0.1 - 10 keV, Effective area of ACIS front (@ 6 keV) ~ 235 cm^2, Energy resolution of HETG (@ 2.5 keV) ~ 5.2 eV, Angular resolution ~ 0.5''.

XMM-Newton: Energy range ~ 0.2 - 12 keV, Effective area of EPIC PN (@ 6 keV) ~ 851 cm^2, Energy resolution of RGS (@ 2.5 keV) ~ 17 eV, Angular resolution ~ 6''.



Suzaku: Energy range ~ 0.4 - 10 keV, Effective area of XIS (@ 6 keV) ~ 1000 cm^2, Energy resolution of XIS (@ 2.5 keV) ~ 80 eV, Angular resolution ~ 1.5'.

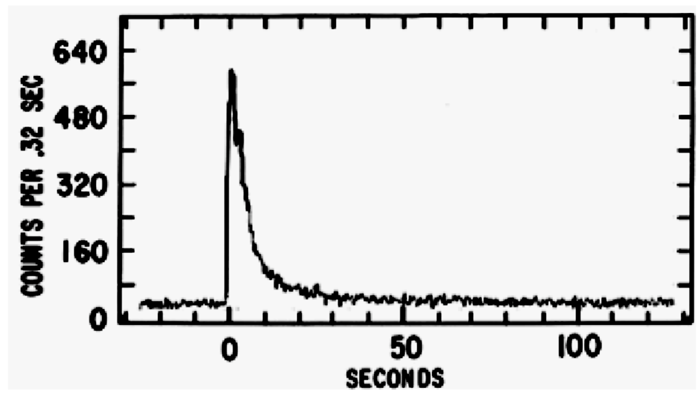
Thermonuclear X-ray Bursts



Accretion on neutron star

Unstable nuclear burning of accreted matter on the neutron star surface causes type I (thermonuclear) X-ray bursts.

Burst light curve



Rise time ≈ 0.5 - 5 seconds
Decay time ≈ 10 - 100 seconds
Recurrence time ≈ hours to day
Energy release in 10 seconds ≈ 10^{39} ergs

↓

Sun takes more than a week to release this energy.

Why is *unstable* burning needed?

Energy release:
Gravitational ≈ 200 MeV / nucleon
Nuclear ≈ 7 MeV / nucleon

Accumulation of accreted matter for hours → Unstable nuclear burning for seconds ⇒ Thermonuclear X-ray burst.

Thermonuclear X-ray Bursts

(1) At T > 8 x 10^7 K, hydrogen burns in a stable manner via hot CNO cycle:

$$^{12}C(p,\gamma)^{13}N(p,\gamma)^{14}O(\beta^+)^{14}N(p,\gamma)^{15}O(\beta^+)^{15}N(p,\alpha)^{12}C$$

(2) Thermal runaway is triggered, that is the burst is ignited, by the temperature sensitivity of helium burning via the 3α reaction:

$$3\alpha \rightarrow {}^{12}C$$

(3) Helium ignition triggers unstable hydrogen burning via the rapid proton capture process (rp-process).

Parameters that set the ignition condition:

(1) chemical composition of accreted matter,
(2) temperature (~ 10^8 K),
(3) column depth (~ 10^8 gm cm^{-2}), and
(4) initial conditions set by the previous bursts.

Why are the thermonuclear X-ray bursts important for understanding neutron stars?

(1) They originate from neutron star surfaces.

(2) Their intensities are ~ 10 times higher than the non-burst emission intensity. This gives higher signal-to-noise ratio.

(3) They show timing and spectral features, that can be used to constrain the *mass*, *radius* and *spin frequency* of the **same** neutron star.

(4) They provide the unique opportunity to understand the thermonuclear flame spreading on neutron star surfaces.

(5) Many bursts are observed from the same neutron star.

(6) Comparatively lower magnetic fields (~ 10^7-10^9 G) of the bursting neutron stars simplify the modeling.

Procedures to constrain neutron star parameters analyzing thermonuclear X-ray bursts:

(1) Spectral studies:
- (a) continuum spectroscopy *(RXTE-PCA),*
- (b) line spectroscopy *(Chandra, XMM-Newton, Suzaku)*.

(2) Studies of fast (millisecond period) timing properties *(RXTE-PCA)*.

Continuum Burst Spectroscopy

★ Burst spectra are normally well fitted with a blackbody model.

★ In principle, neutron star radius can be measured from the observed bolometric flux (F_{obs}) and blackbody temperature (T_{obs}), and the known source distance (d):

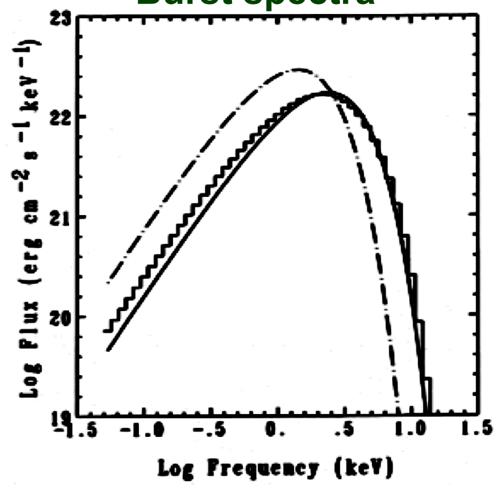
$$R_{obs}=d.(F_{obs}/(\sigma T^4_{obs}))^{1/2}$$

★ But there are systematic uncertainties:
(1) unknown amount of spectral hardening due to electron scattering;
(2) effect of unknown gravitational redshift.

$$T = T_{obs}.(1+z)/f \qquad R = R_{obs}.f^2/(1+z)$$

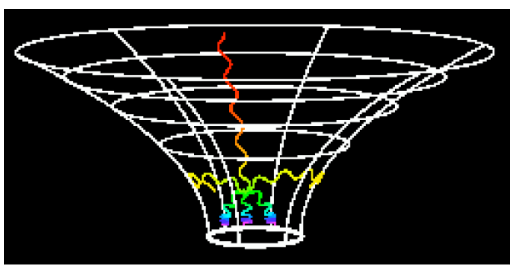
$$\begin{cases} z > 0;\ f \sim 1.0 - 2.0 \\ 1+z = [1 - (2GM/Rc^2)]^{-1/2} \end{cases}$$

Burst spectra



London, Taam & Howard (1986)

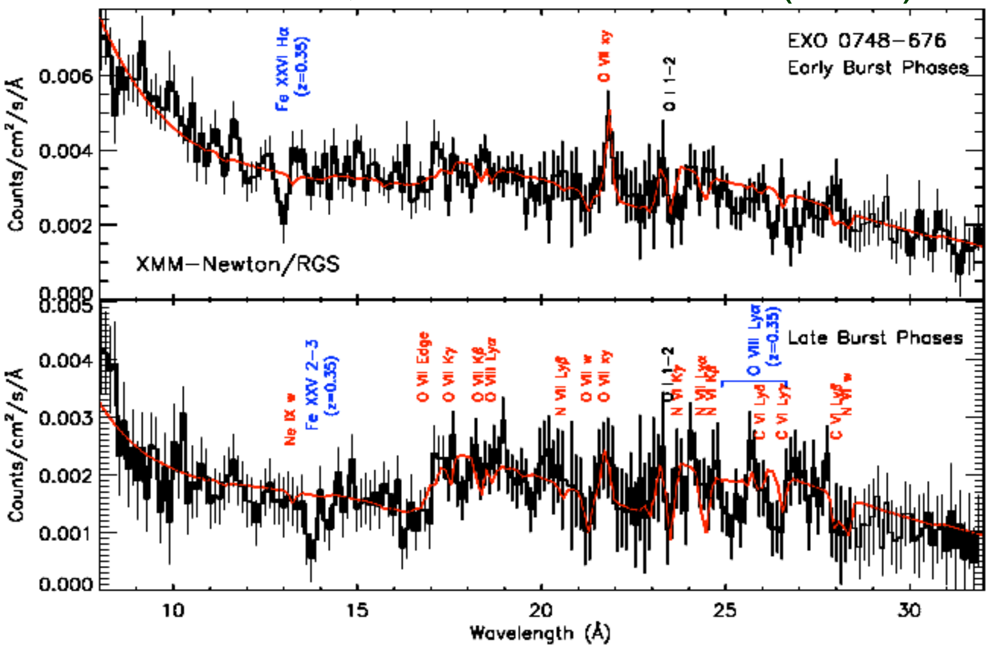
Gravitational redshift



Chemical composition of neutron star atmosphere ⇒ f
Neutron star radius-to-mass ratio ⇒ 1+z

Line Burst Spectroscopy

Cottam, Paerels & Mendez (2002)



XMM-Newton grating observations of surface atomic spectral absorption lines during X-ray bursts from an LMXB (EXO 0748-676): measured gravitational redshift $1+z = 1.35$, and hence $Rc^2/GM = 4.4$.

⇓

Observation of surface atomic spectral line at the energy E_{obs}

⇓

Identification: original line energy = E_0

⇓

Gravitational redshift $1+z = E_0/E_{obs}$

⇓

Neutron star "radius to mass" ratio from $1+z = [1-(2GM/Rc^2)]^{-1/2}$

But why LMXBs and X-ray bursts?

These Fe absorption lines could be produced in the upper atmosphere of the neutron star, and the continuous accretion might supply the Fe ions.

Line Burst Spectroscopy

Why LMXBs and X-ray bursts?

* Comparatively lower magnetic field (10^7-10^9 G):
  (1) magnetic splitting is negligible: line identification is easier;
  (2) magnetic field does not complicate the modeling of neutron star atmosphere and photon emission.

* For isolated neutron stars, heavy elements do not exist in the atmosphere. For LMXBs and during bursts, continuous accretion (at a low rate) and radiative pressure may keep heavy elements in the atmosphere for the time required for spectral line detection.

* During the bursts, high photon flux from the neutron star surface provides good signal-to-noise ratio.

Line Burst Spectroscopy

- But the neutron stars in LMXBs normally spin very fast due to accretion induced angular momentum transfer.

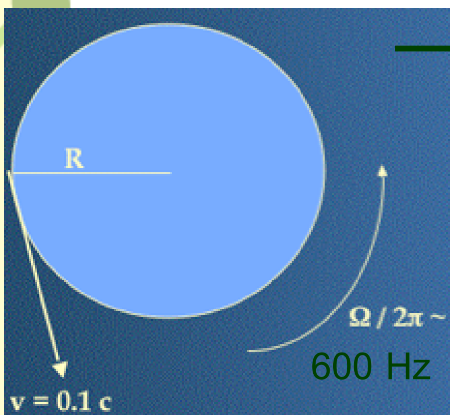


Figure courtesy F. Ozel.

→ *Spinning neutron star:* surface speed is ~ 0.1c; Doppler effect will make the spectral line broad and asymmetric.

How do we measure (1+z) from a broad and skewed line?

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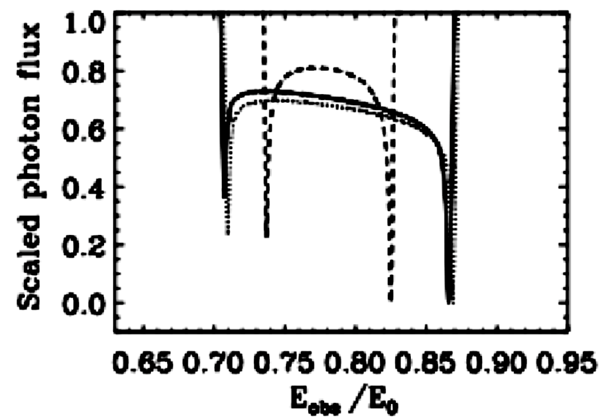
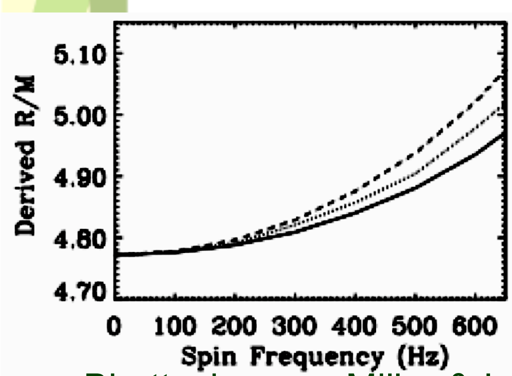
$$\begin{cases} E_{gm} = (E_1 E_2)^{1/2} \\ 1+z = E_0/E_{gm} \\ Rc^2/GM = 2.(1 - (1+z)^{-2})^{-1} \end{cases}$$

⇓



Bhattacharyya, Miller & Lamb (2006)



Bhattacharyya, Miller & Lamb (2006)

← Better than 2% estimate!

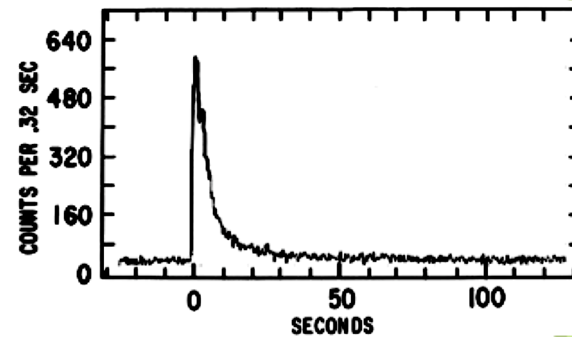
Modeling of the shapes of the spectral lines will be useful to constrain other neutron star parameters.

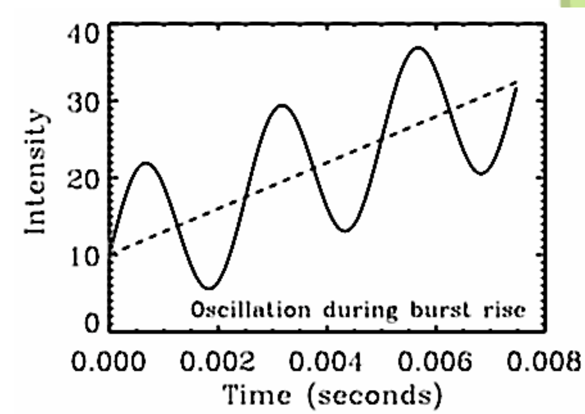
Fast Timing Properties of X-ray Bursts (Burst Oscillations)

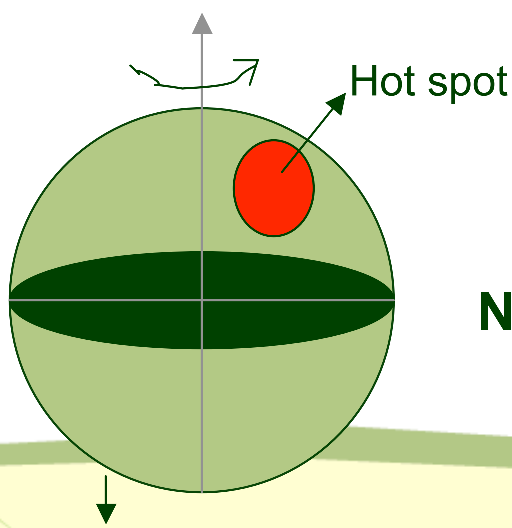
- What are burst oscillations?
  These are millisecond period variations of observed intensity during thermonuclear X-ray bursts.

- What is their origin?
  Asymmetric brightness pattern on the spinning neutron star surfaces.

Burst light curve







Neutron star spin frequency = Burst oscillation frequency

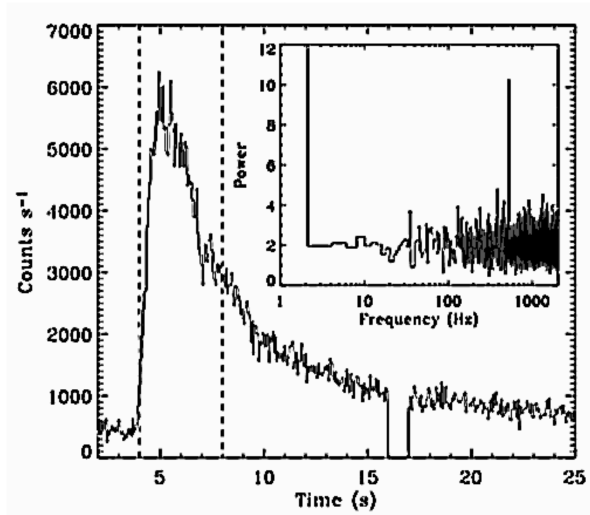
Burst Oscillations

Burst oscillation sources:

Number	LMXB	Spin frequency (Hz)
1	EXO 0748–676	45
2	4U 1916-05	270
3	XTE J1814–338[P]	314
4	4U 1702–429	330
5	4U 1728–34	363
6	SAX J1808.4–3658[P]	401
7	SAX J1748.9–2021	410
8	KS 1731–260	524
9	1A 1744–361	530
10	Aql X–1	549
11	X1658–298	567
12	4U 1636–53	582
13	X1743–29	589
14	SAX J1750.8–2980	601
15	4U 1608–52	619

The spin frequencies of neutron stars in 13 LMXBs are known from burst oscillations.

Bhattacharyya et al. 2006



Measured spin rates can be used to constrain neutron star mass and radius.

RXTE-PCA is the only current instrument that can detect burst oscillations. In future, Indian instrument ASTROSAT-LAXPC will be able to do it.

Burst Oscillations: Stellar Mass and Radius

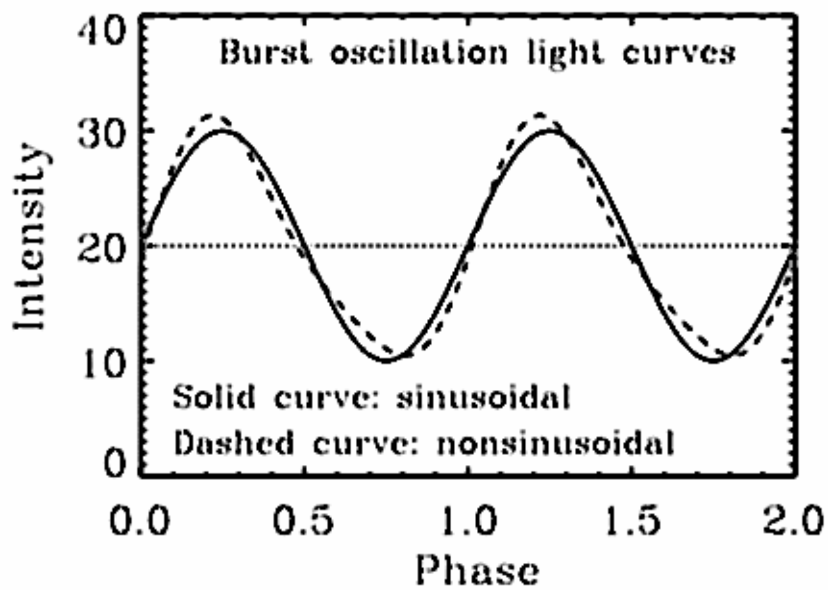
- Modeling of burst oscillation amplitudes and light-curve-shapes :

  $\Downarrow$

  Neutron star mass and radius-to-mass ratio.

- Models should include the following physical effects: Doppler effect, special relativistic beaming, gravitational redshift, gravitational light bending, frame dragging, etc.

- However non-sinusoidal burst oscillation light curves are required to fully utilize this procedure.



Modeling Burst Oscillation Light Curves

Non-sinusoidal light curves from the decay portions of the X-ray bursts from the LMXB XTE J1814-338.

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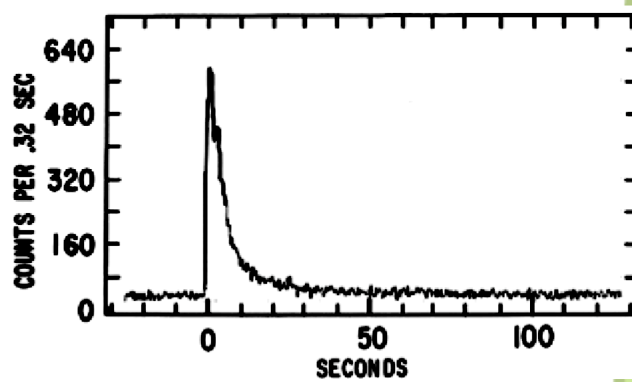
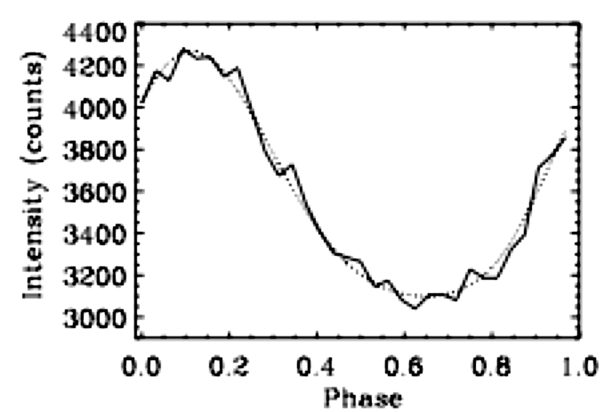
Fitting the observed burst oscillation light curves with our theoretical model (assuming a hot spot on the spinning neutron star surface), we have constrained a few parameters, including stellar radius-to-mass ratio.

Burst light curve



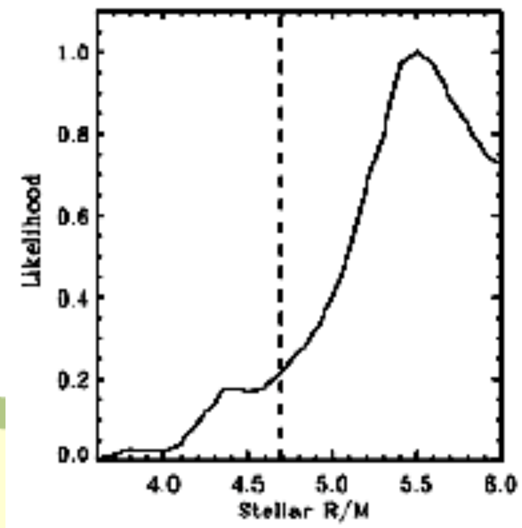
XTE J1814-338
(RXTE-PCA data)



Bhattacharyya et al. (2005)

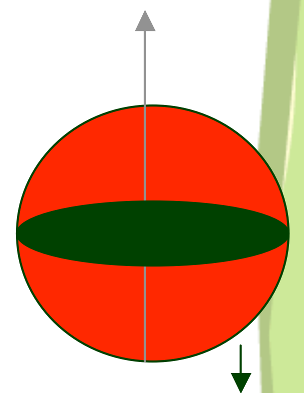
The vertical dashed line gives the lower limit of the stellar radius-to-mass ratio with 90% confidence.

XTE J1814-338
(RXTE-PCA data)



But what causes oscillations during the burst decay?

What causes the brightness (and temperature) asymmetry on the neutron star surface?



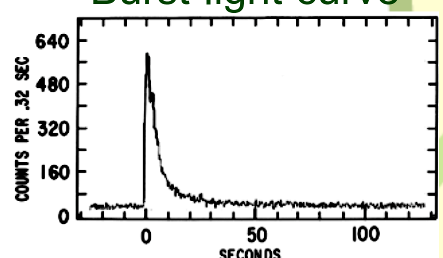
Spinning neutron star

Possibilities:

(1) Hot spot created by vortex or surface magnetic field.

(2) Surface waves in the burning layer or photosphere in the equatorial region may cause temperature variation on the stellar surface.

Burst light curve



However, it is unlikely for surface waves to produce the significantly non-sinusoidal burst oscillation light curves observed from XTE J1814-338.

Burst Oscillations during the Burst Rise

Hot spot ⇒ Oscillations during the burst rise.

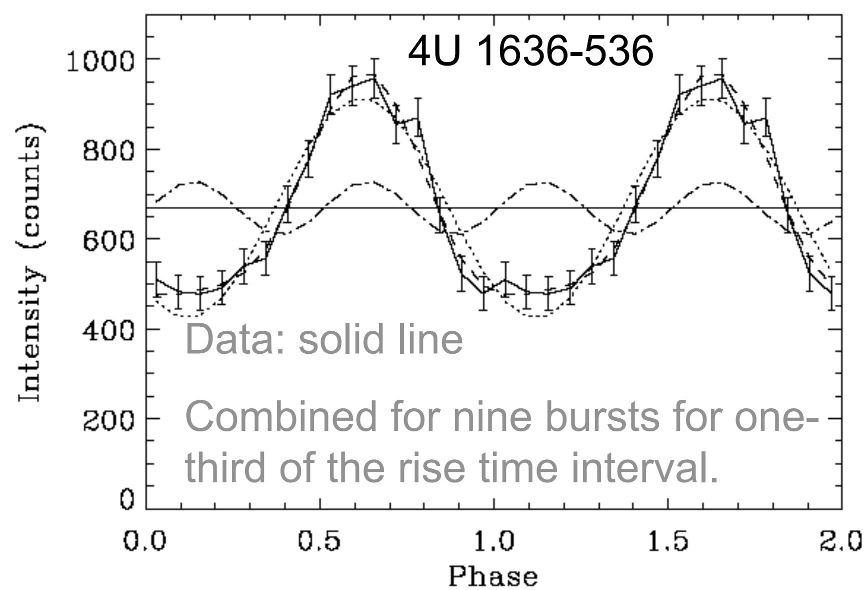
Reasons:

(1) The expanding hot spot (burning region) can naturally give rise to oscillations during burst rise.

(2) The observed amplitude of oscillations during the early burst rise is large, which is expected from a small hot spot.

(3) A small hot spot (burning region) expected during early burst rise can give rise to significantly *non-sinusoidal* burst oscillation light curve. **We have recently, for the first time, discovered such a light curve from the early rising phase of the bursts from an LMXB.**

Bhattacharyya & Strohmayer 2005
(*RXTE-PCA data*)



Burst rise oscillations are promising for constraining neutron star mass and radius-to-mass ratio.

Problem:

Early phase of burst rise exist for less than a second, and the total number of detected photons is comparatively smaller. Hence, for many bursts, signal-to-noise ratio is not high.

Nevertheless, it is worthwhile to study the burst rise oscillations. Such a study will also provide the opportunity to understand the thermonuclear flame spreading on the neutron star surfaces.

Summary of constraining neutron star EOS models

EOS ⇐ spin, mass and radius of the **same** neutron star.

Thermonuclear X-ray bursts give the opportunity of three types of studies: *continuum spectroscopy*, *line spectroscopy* and *fast timing study*.

Burst oscillations ⇒ Neutron star **spin frequency**

Surface atomic spectral line or burst oscillations ⇒ stellar **Rc^2/GM**
Study of bursts and accretion flow ⇒ chemical composition of stellar atmosphere

⇓

Continuum spectroscopy ⇒ Stellar **radius**

<u>Example</u>: LMXB EXO 0748-676:
Spin frequency = 45 Hz (burst oscillations)
Rc^2/GM = 4.4 (line spectroscopy)
R or M = ?

Thermonuclear Flame Spreading on Neutron Stars

When does it happen?

During the thermonuclear X-ray bursts (mostly during burst rise).

Why should we care?

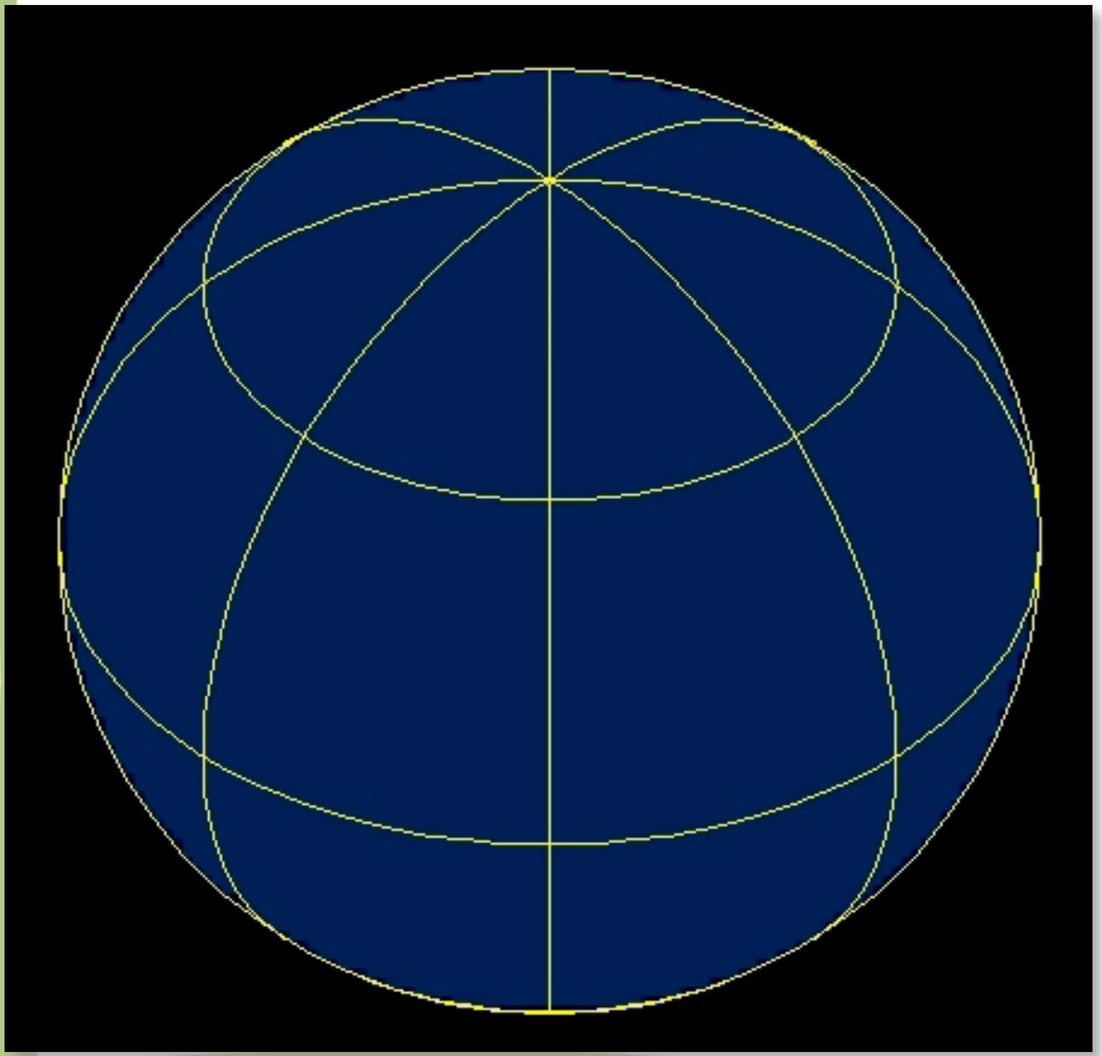
(1) It is an interesting research field on its own. It is basically atmospheric physics under extreme conditions: extreme gravity, high density (10^5-10^6 gm/cc), high magnetic field, huge energy generation and radiation pressure, large stellar spin (and hence Coriolis force), etc.

(2) It can be useful to understand the neutron star atmosphere, and to constrain surface magnetic field, chemical composition of matter, etc. It is also useful to model burst rise oscillations.

Theoretical study:

Not yet done taking all the main physical effects into account. Until recently, observations could not provide enough motivation. Our recent observational findings may provide this motivation.

Thermonuclear Flame Spreading on Neutron Stars

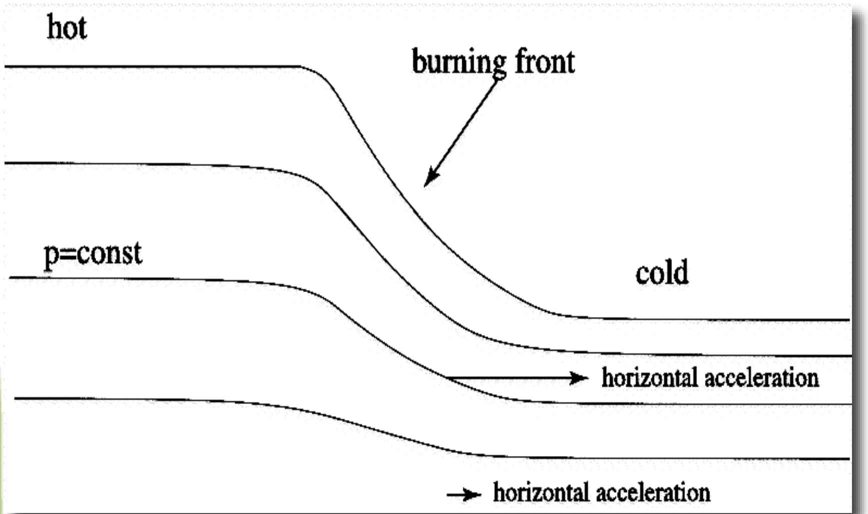


A simulation considering Coriolis force, but ignoring several other physical effects (such as surface magnetic field).

Thanks to Anatoly Spitkovsky!

Theoretical modeling of thermonuclear flame spreading

Spitkovsky et al. (2002)



Neutron star spin frequency 300-600 Hz ⇒ Coriolis force important.

Thin burning layer ⇒ Geostrophic approximation.

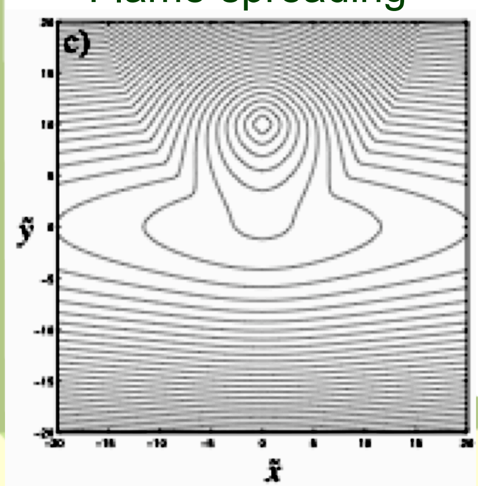
Flame speed ~ Ageostrophic speed.

For weak turbulent viscosity, flame speed ϑ ~ 5 - 20 km/s.

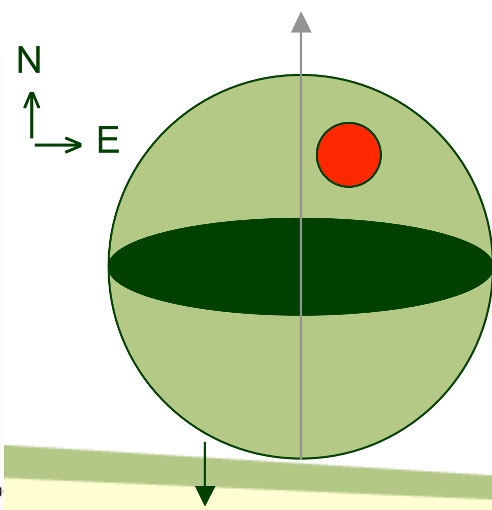
For strong turbulent viscosity, flame speed $\vartheta \leq$ 300 km/s.

$\vartheta_{pole} < \vartheta_{equator}$

Flame spreading



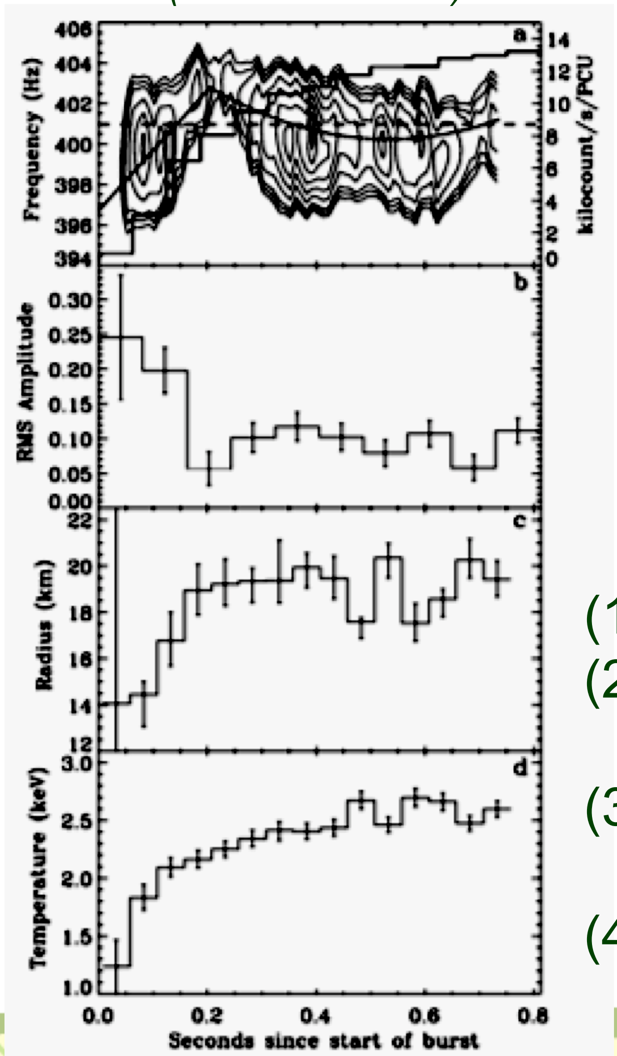
Spitkovsky et al. (2002)



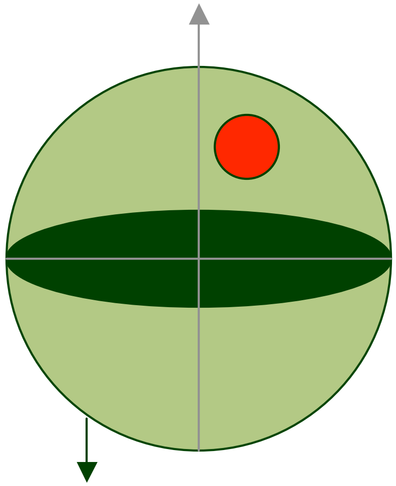
Spinning neutron star

Thermonuclear Flame Spreading on Neutron Stars

SAX J1808.4-3658
(RXTE-PCA data)

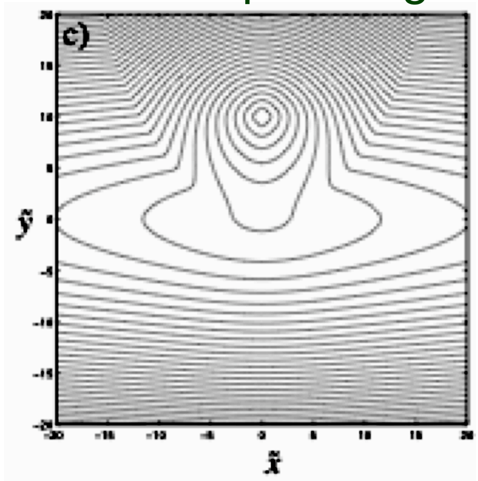


Bhattacharyya & Strohmayer (2006c)



Spinning neutron star

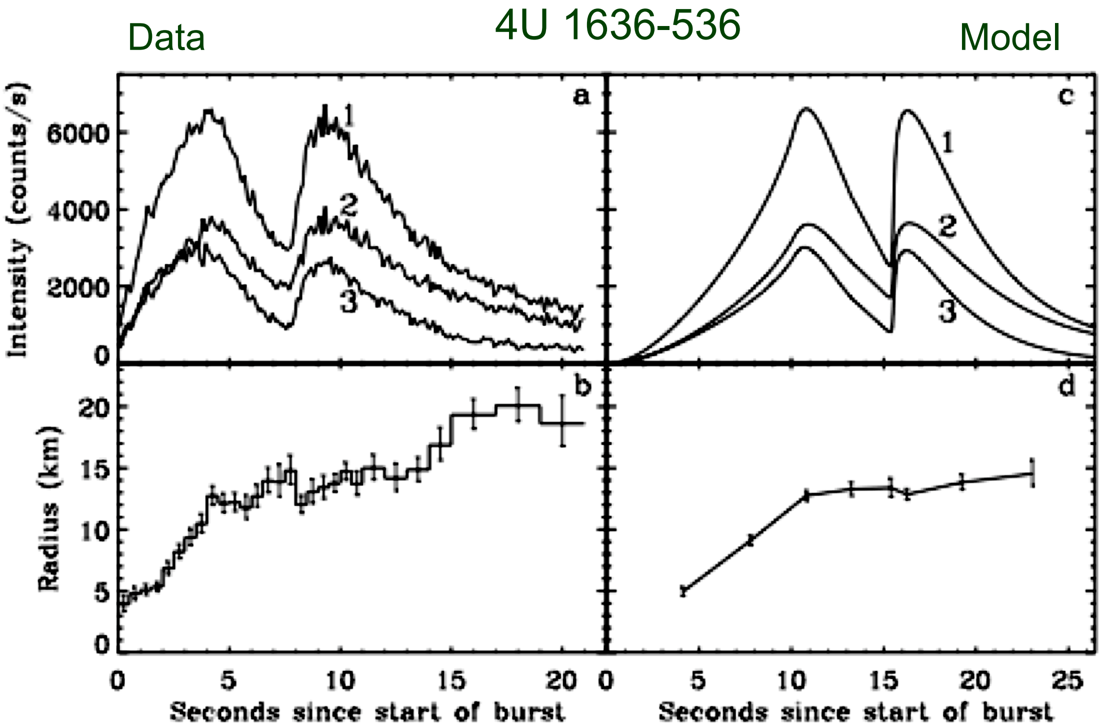
Flame spreading

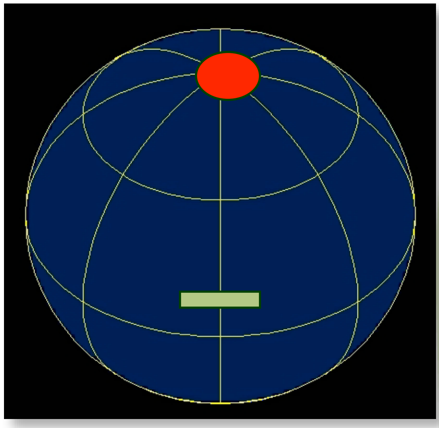


Spitkovsky et al. (2002)

(1) Initial large amplitude is due to small hot spot.
(2) As the burning region grows, amplitude decreases and radius increases quickly.
(3) The low amplitude after 0.2 second is due to the residual asymmetry.
(4) The complex frequency evolution may be due to the slow eastward or westward acceleration of the center of the burning region. Magnetic field may become dynamically important.

Thermonuclear Flame Spreading on Neutron Stars: Weak Double-peaked X-ray Bursts





Neutron star with polar ignition

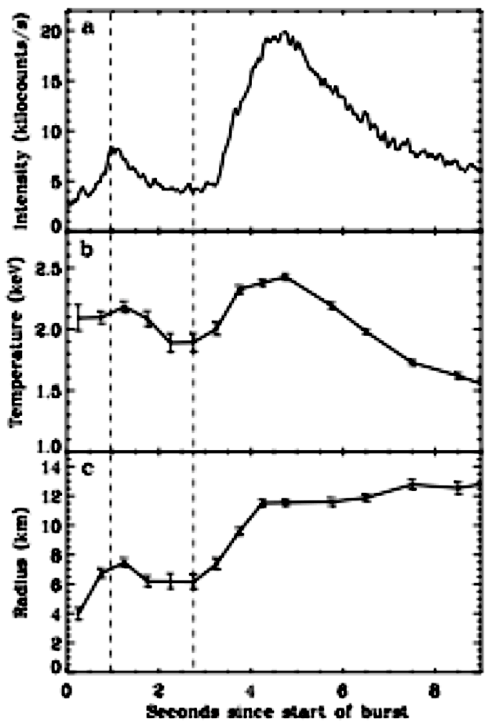
Bhattacharyya & Strohmayer (2006a); *RXTE-PCA data*

(1) Burst ignition at a pole, which explains the lack of oscillations and the rarity of the burst.
(2) Azimuthally symmetric temporary burning front stalling cools the burning region for a few seconds, while keeping the burning area unchanged. This can explain the intensity and temperature drop during the dip.
(3) The subsequent expansion of burning region explains the second intensity peak.

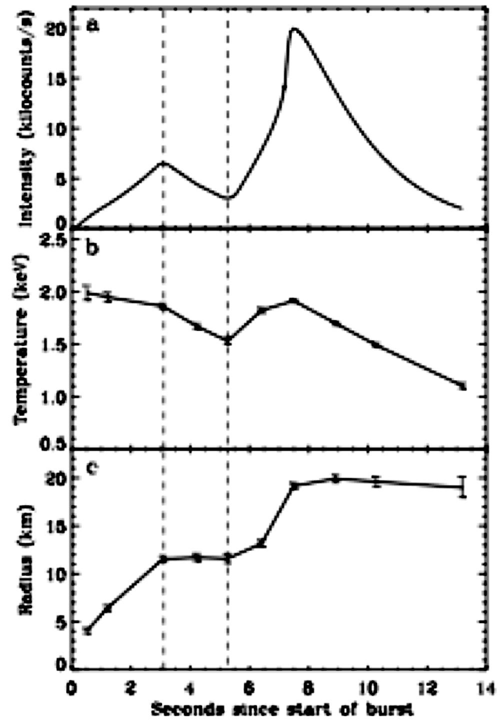
Thermonuclear Flame Spreading on Neutron Stars: Weak Double-peaked X-ray Bursts

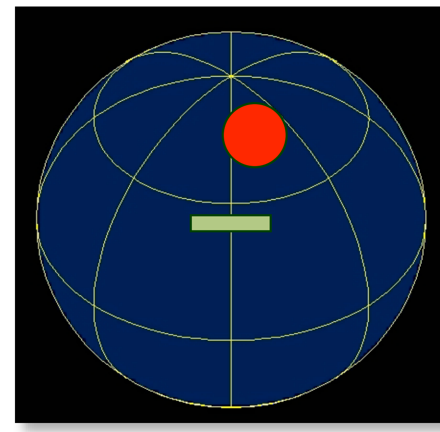
4U 1636-536

Data



Model



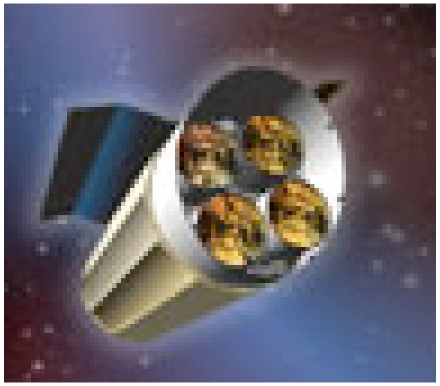


Neutron star with high latitude ignition

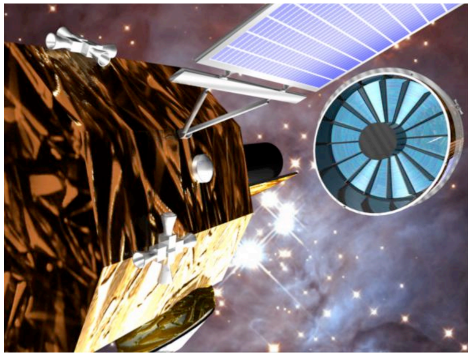
Bhattacharyya & Strohmayer (2006b); *RXTE-PCA data*

Vertical dashed lines give the time interval in which the radius (and hence the burning region area) does not change much and the temporary burning front stalling occurs.

Some Future X-ray Missions



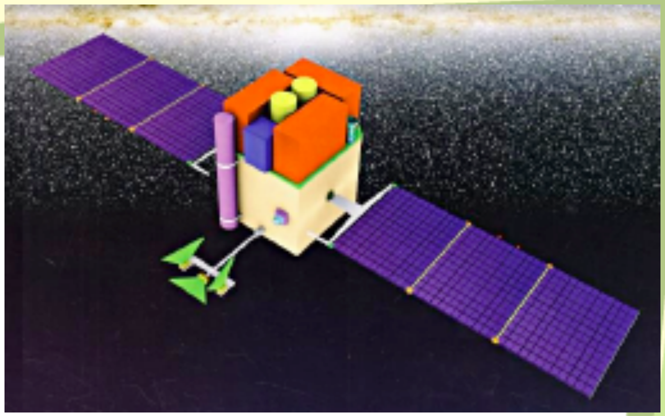
Constellation-X
(NASA)



The X-ray Evolving Universe Spectrometer (XEUS)
(ESA)

ASTROSAT

(India's proposed multiwavelength astronomy mission)



Instruments:

LAXPC: Continuum burst spectroscopy, burst oscillations and thermonuclear flame spreading. **This instrument will directly measure neutron star spin frequencies.**

SSM: Detection of transient LMXBs during outbursts. Many transient LMXBs exhibit thermonuclear X-ray bursts.

SXT, CZTI, UVIT: Accretion flow: continuum and line spectroscopy, and simultaneous X-ray, UV and optical observations $\Rightarrow$ chemical composition of accreted matter, accretion rate, etc.

Conclusions

- Studies of thermonuclear X-ray bursts can be very useful to constrain the spin rate, mass and radius of the **same** neutron star ⇒ **EOS model** of high density cold matter in the neutron star cores.

- Extensive observation and analysis of the data from the rising portions of the bursts ⇒ modeling of burst oscillations and thermonuclear flame spreading.

- Theoretical study of thermonuclear flame spreading on the rapidly spinning neutron stars should be done considering all the main physical effects (including magnetic field, nuclear energy generation, Coriolis effect, strong gravity, etc.).

❊❊❊ *Thank you!* ❊❊❊