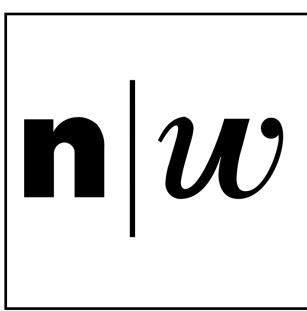


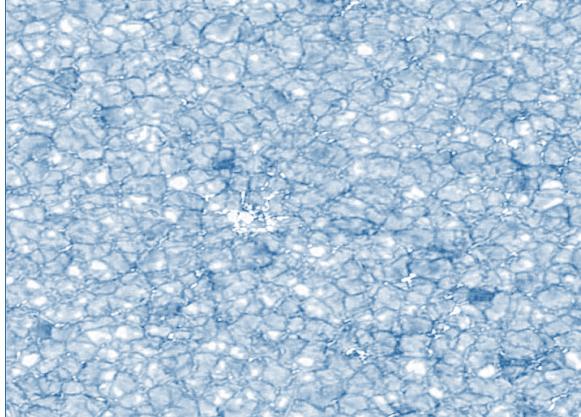


Predicting solar flares with IRIS spectra and machine learning

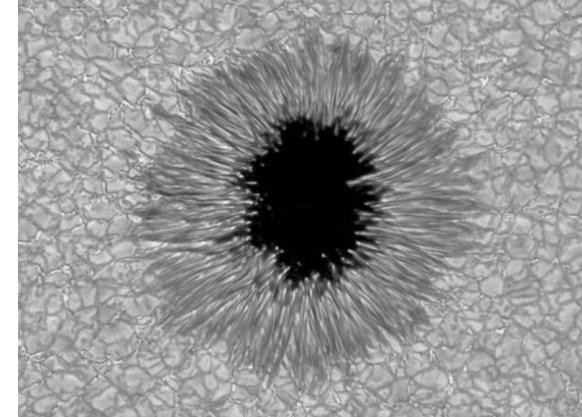


New data source (IRIS Mg II Spectra)

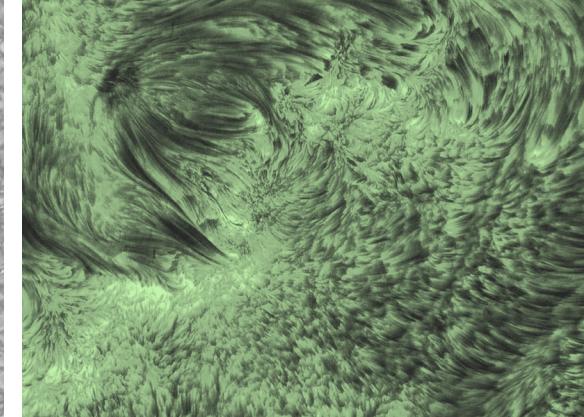
Quiet Sun



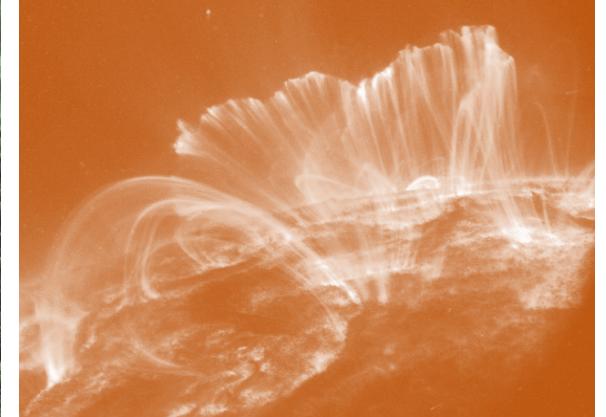
Sunspot



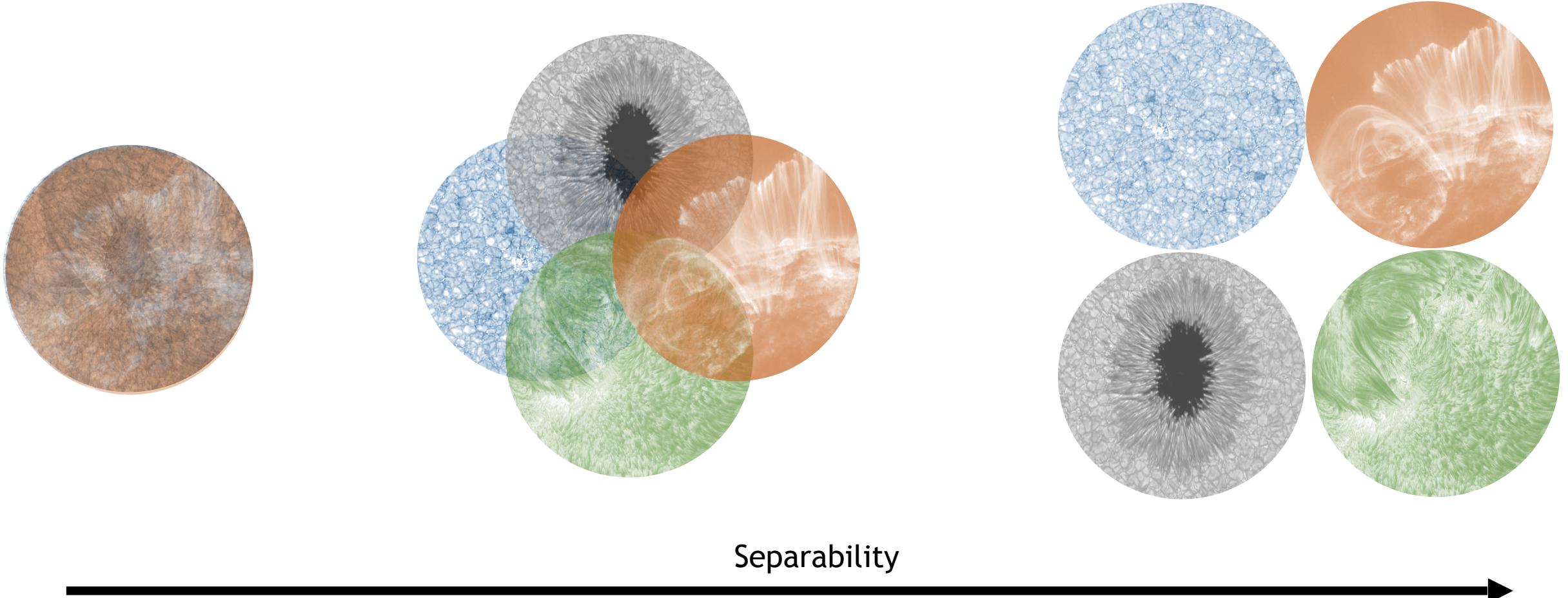
Active region



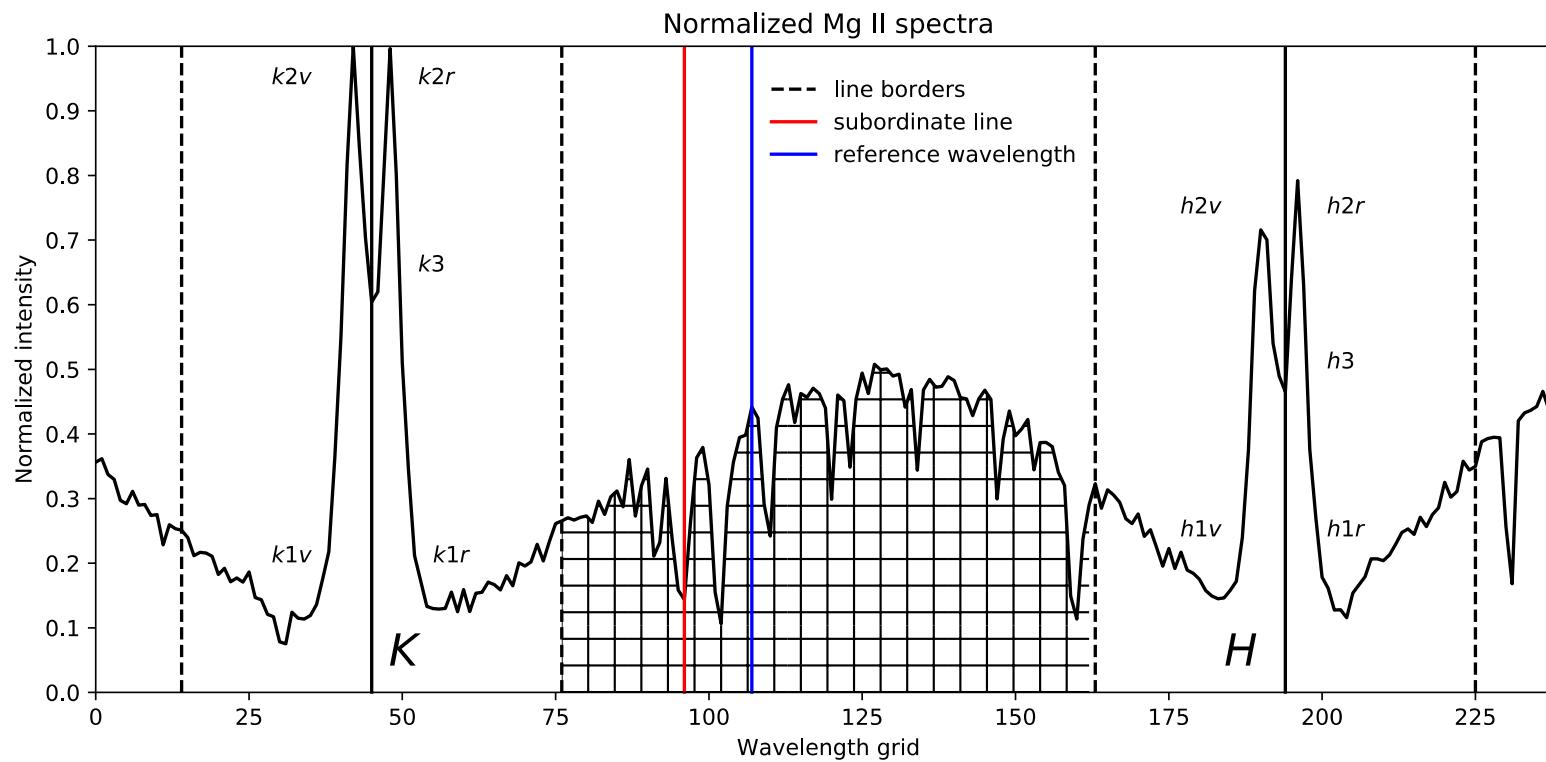
Pre-Flare



What do we plan on doing with these four data sets?

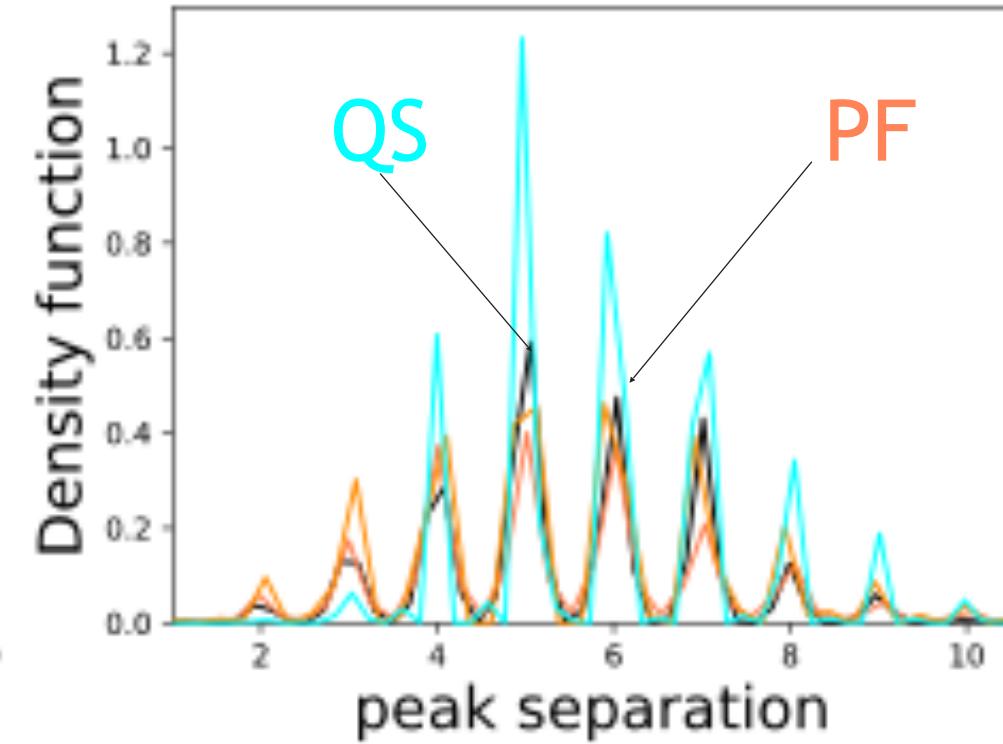
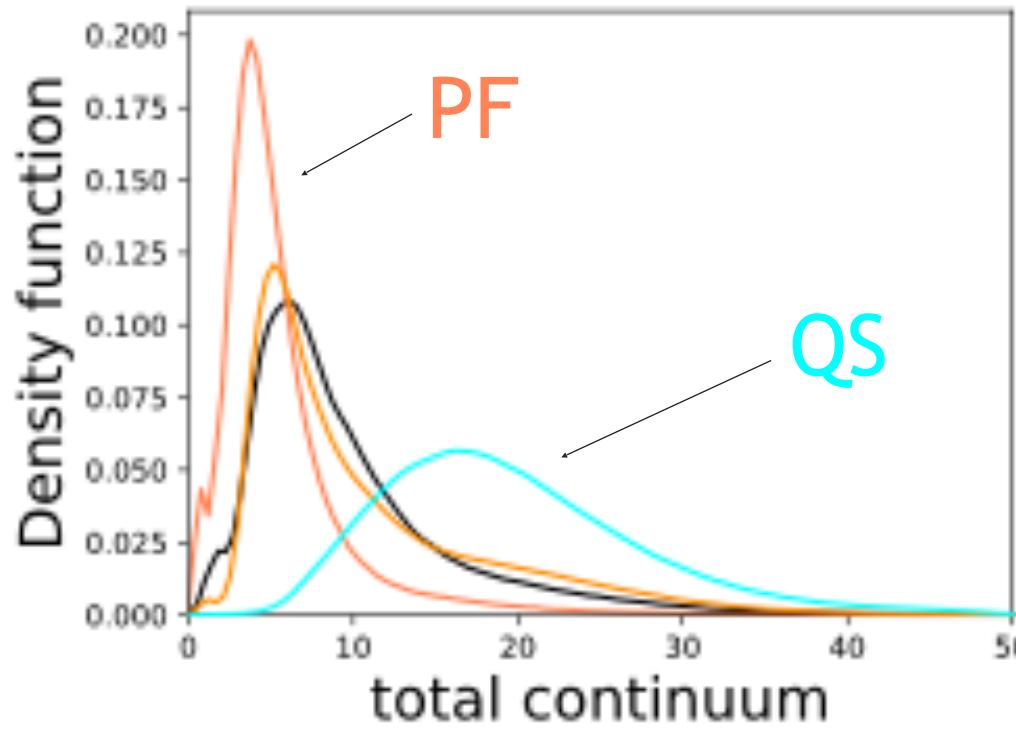


Informative basis



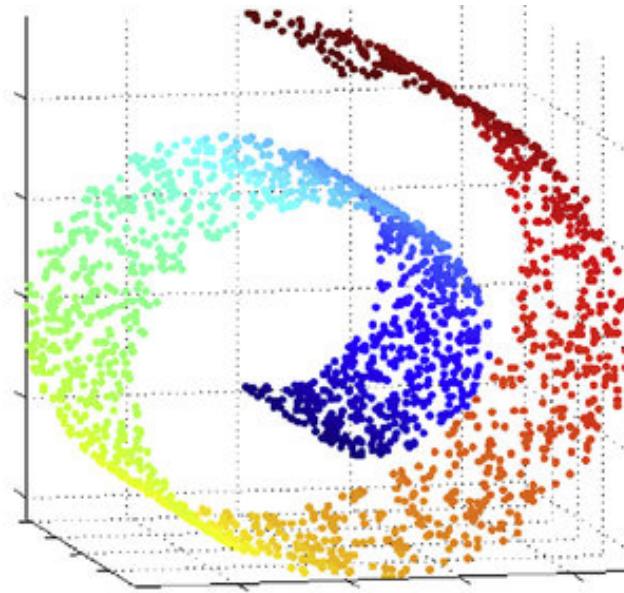
1. k/h ratio
2. Intensity
3. K3-height
4. Peak ratio
5. Line width
6. Line center
7. Triplet emission
8. Peak separation
9. Line asymmetry
10. Continuum emission

Feature distributions of the 4 data sets

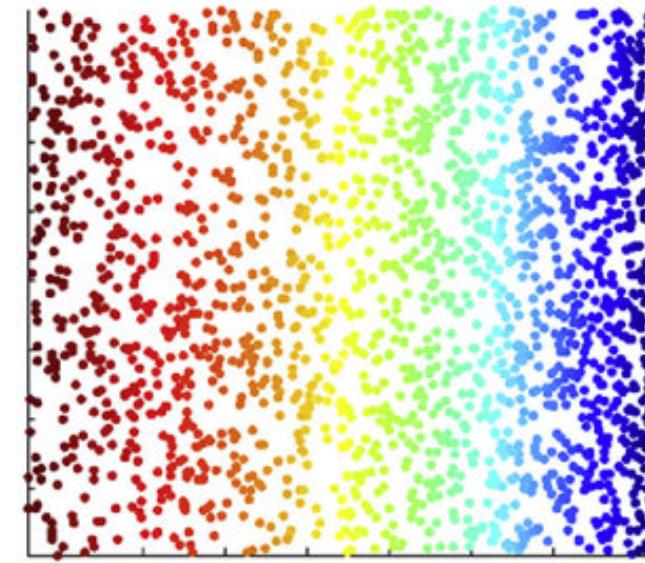


- C0 (distributions)
- C1 (PCA)
- C2 (t-SNE)
- C3 (neural networks)

Dimensionality reduction



High (3d)

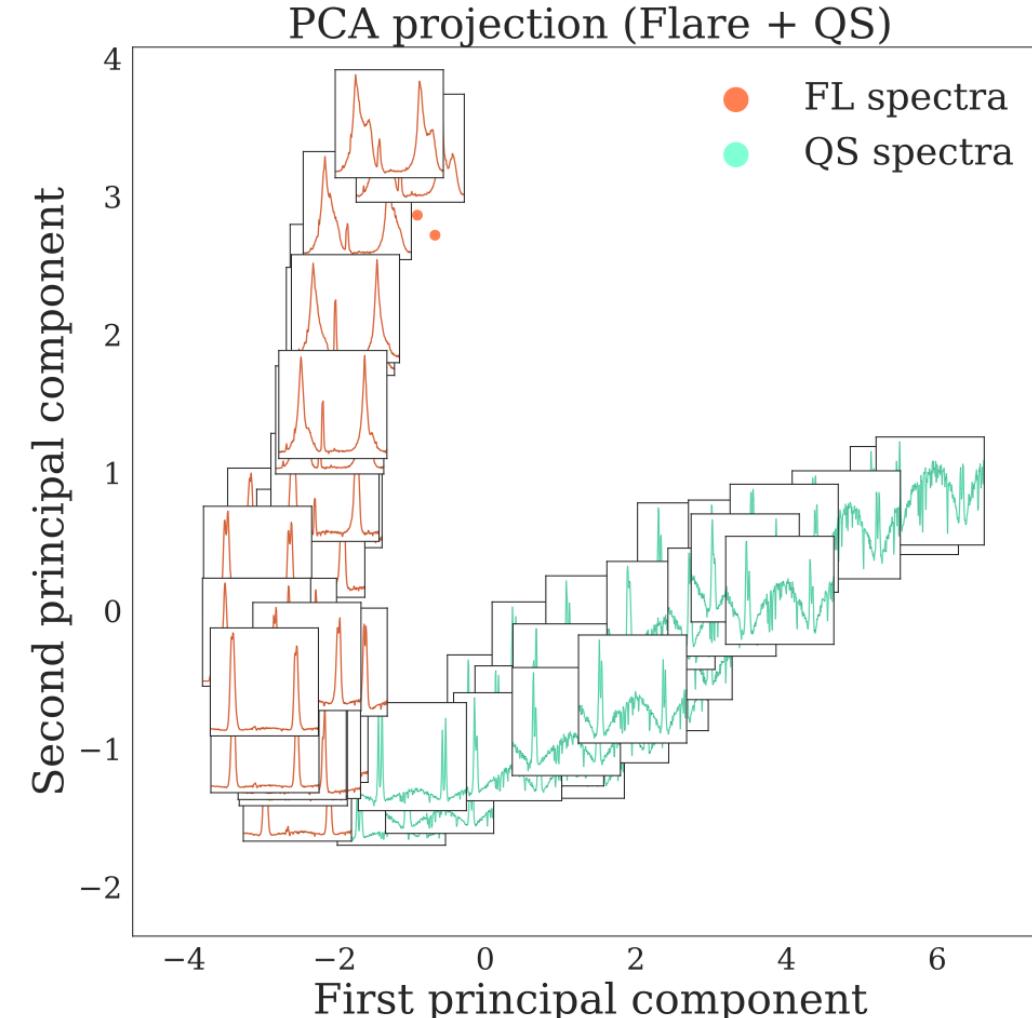


Low (2d)

n|w

PCA

C1

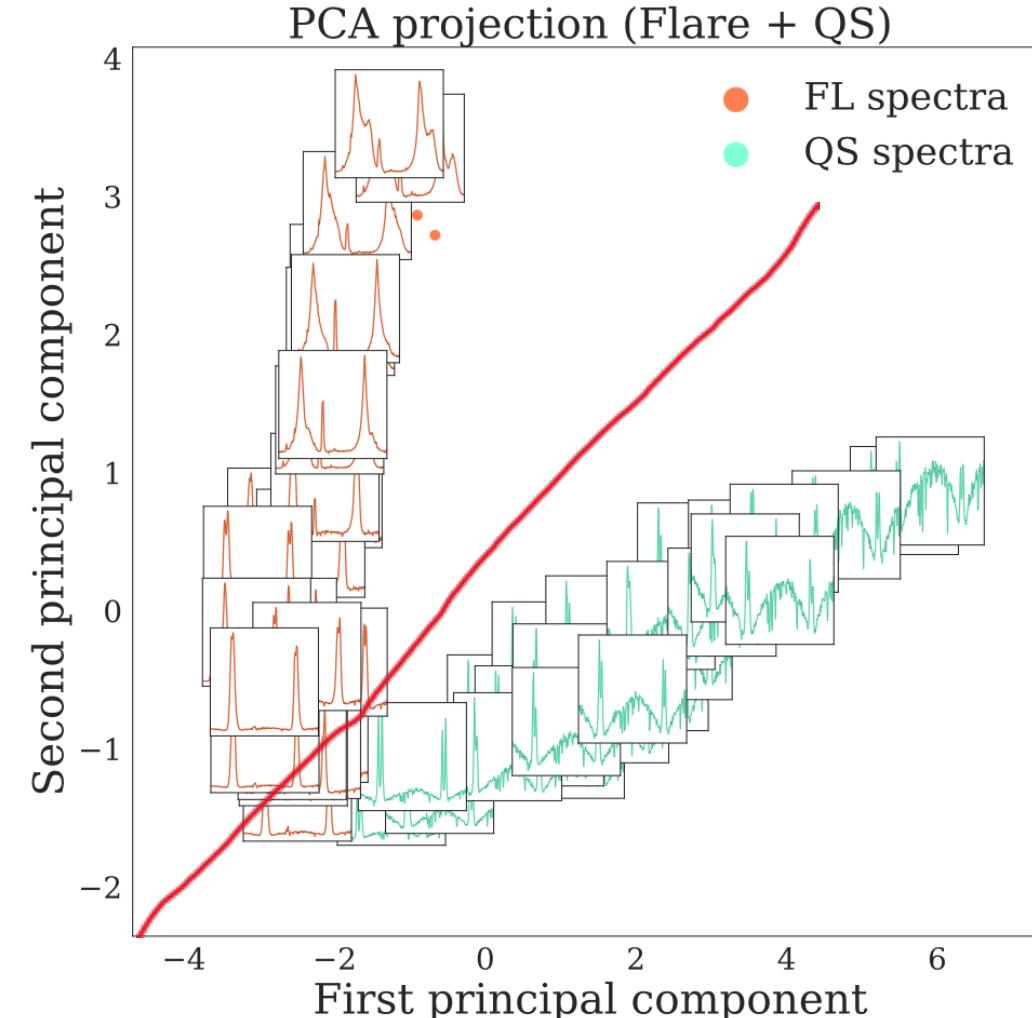


- C0 (distributions)
- C1 (PCA)
- C2 (t-SNE)
- C3 (neural networks)

n|w

PCA

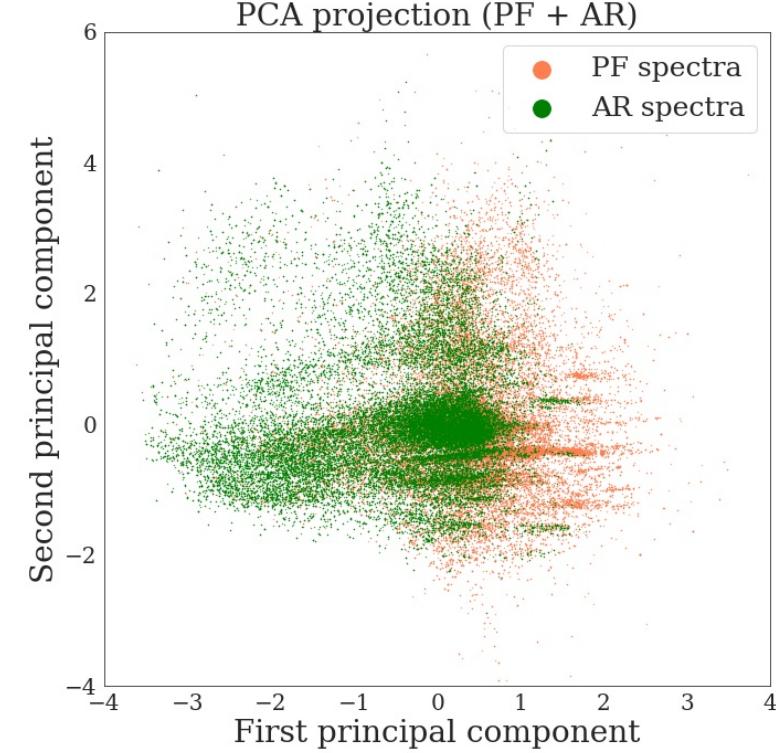
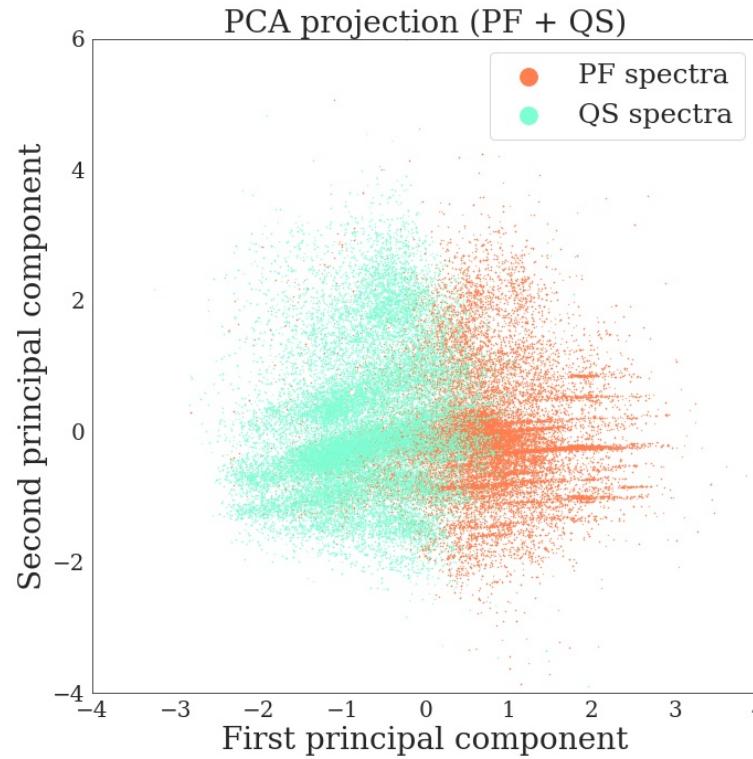
C1



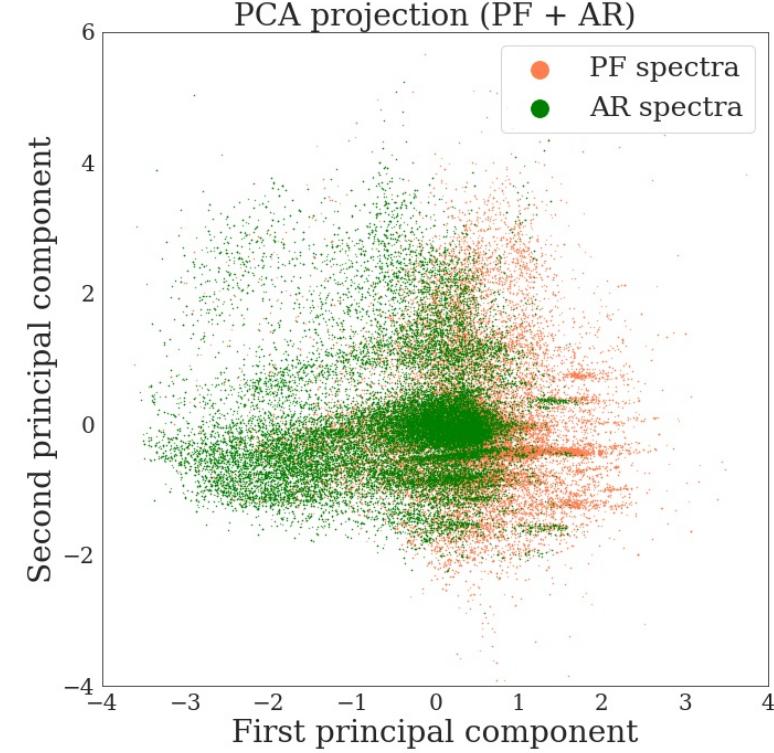
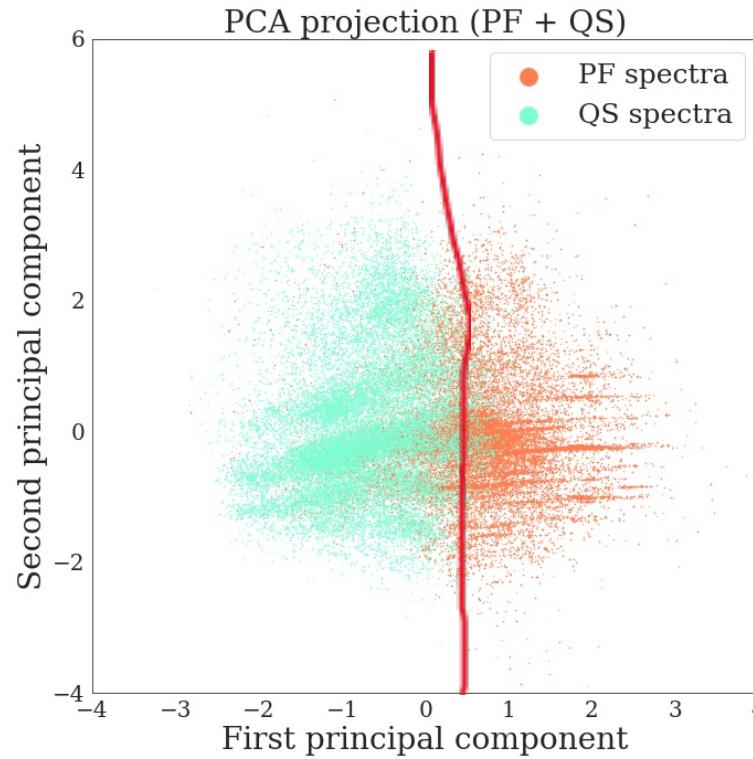
- C0 (distributions)
- C1 (PCA)
- C2 (t-SNE)
- C3 (neural networks)

PCA

C1



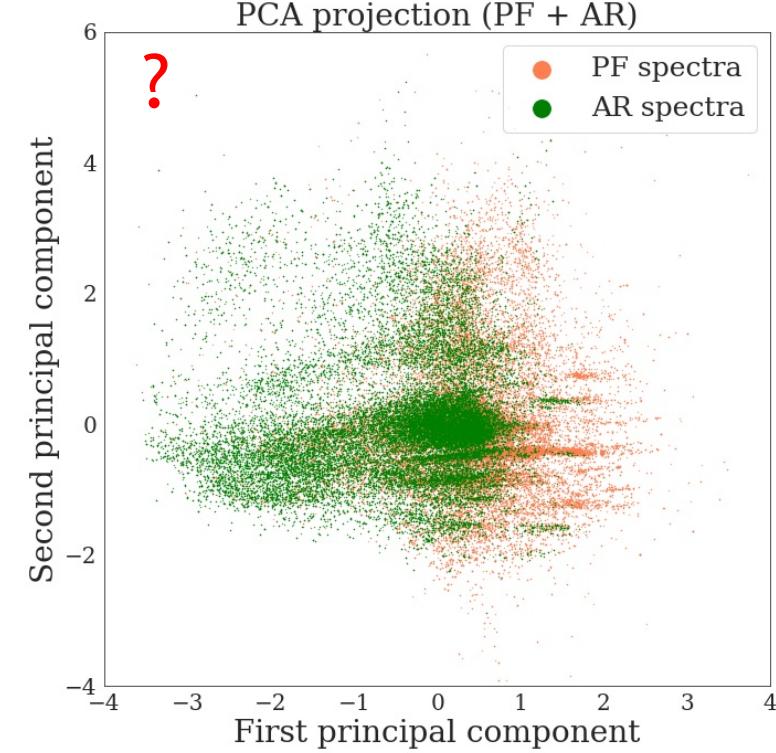
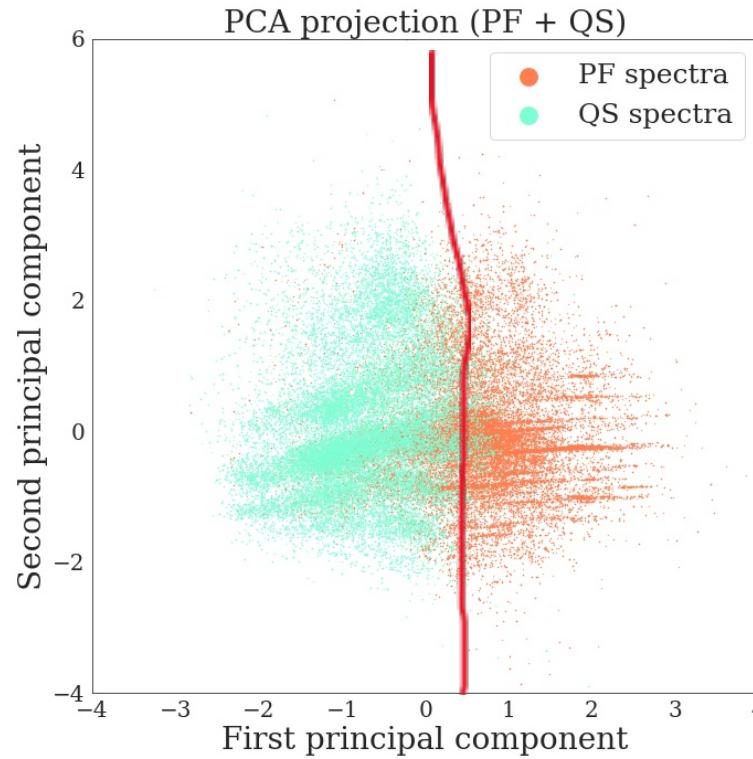
- C0 (distributions)
- C1 (PCA)
- C2 (t-SNE)
- C3 (neural networks)



- C0 (distributions)
- C1 (PCA)
- C2 (t-SNE)
- C3 (neural networks)

PCA

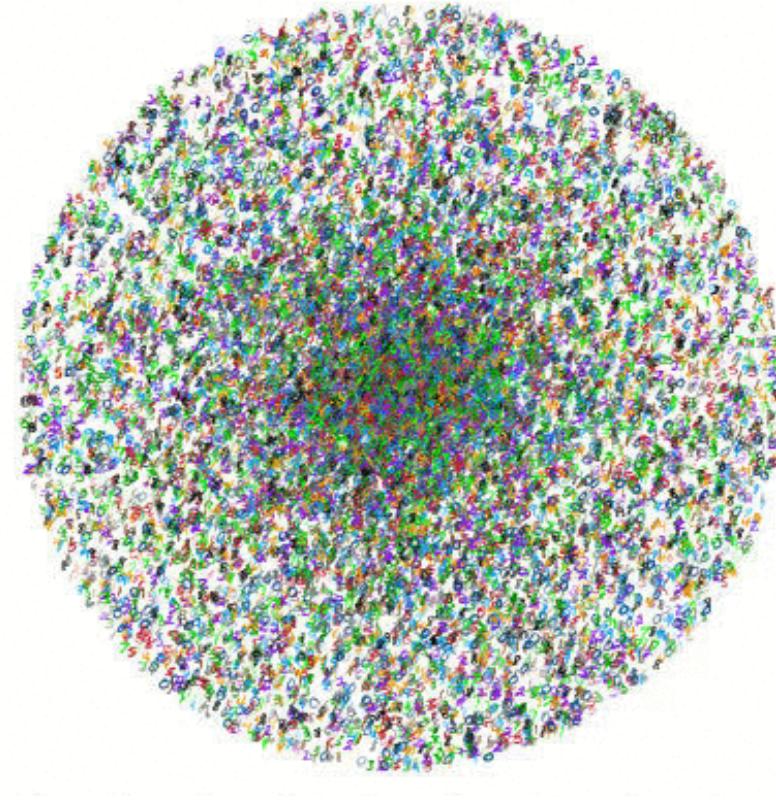
C1



- C0 (distributions)
- C1 (PCA)
- C2 (t-SNE)
- C3 (neural networks)

t-SNE

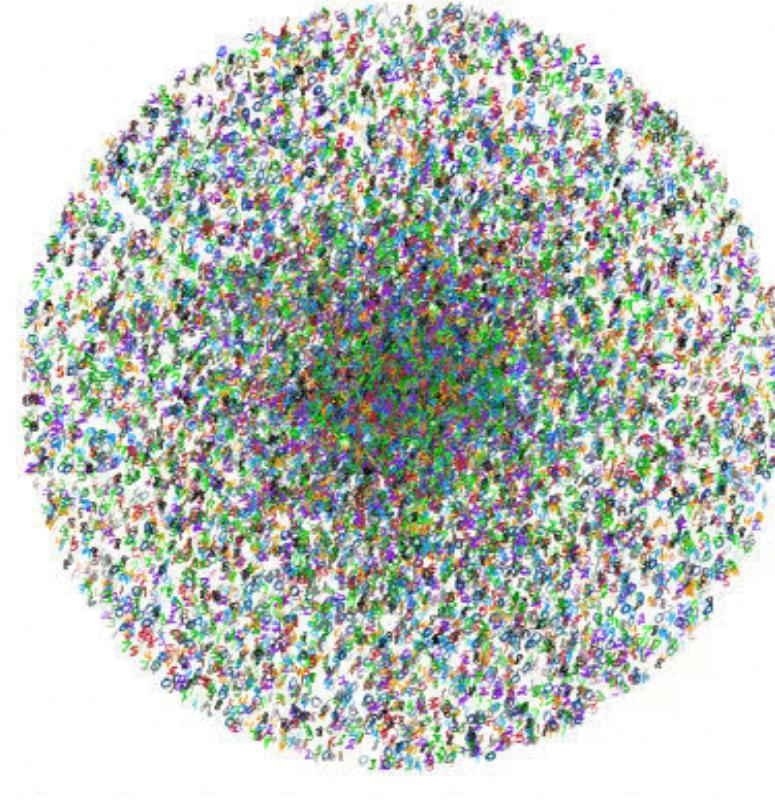
C2



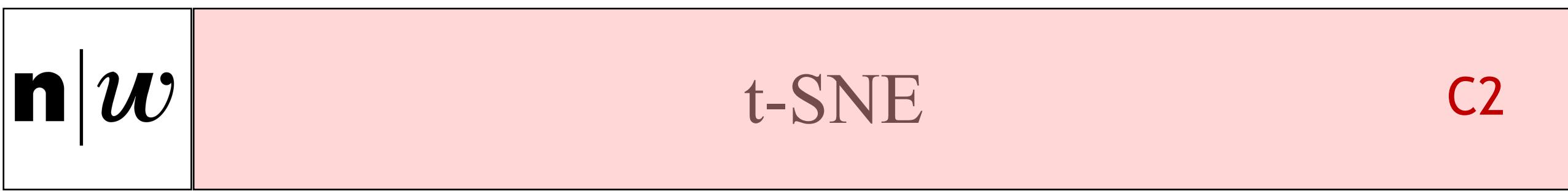
A t-SNE visualization of the IRIS dataset.

- C0 (distributions)
- C1 (PCA)
- C2 (t-SNE)
- C3 (neural networks)

t-SNE



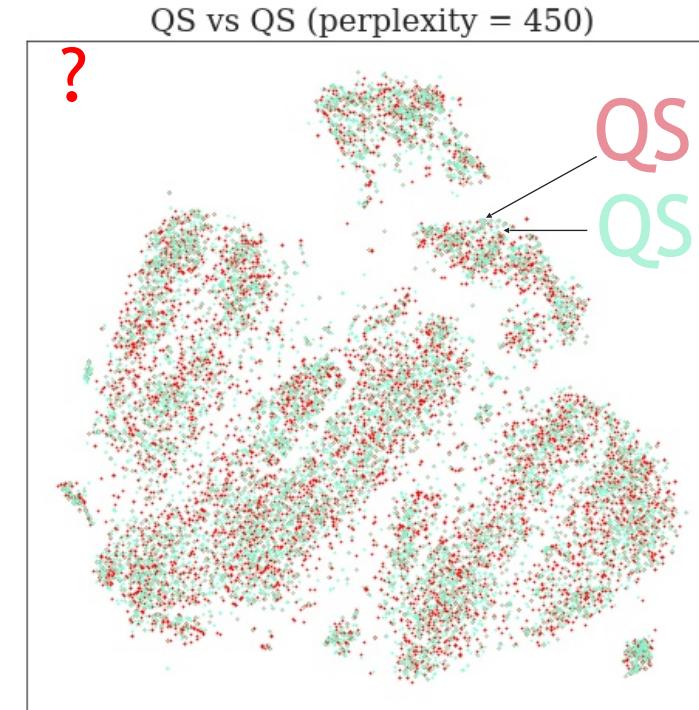
- C0 (distributions)
- C1 (PCA)
- C2 (t-SNE)
- C3 (neural networks)



n|w

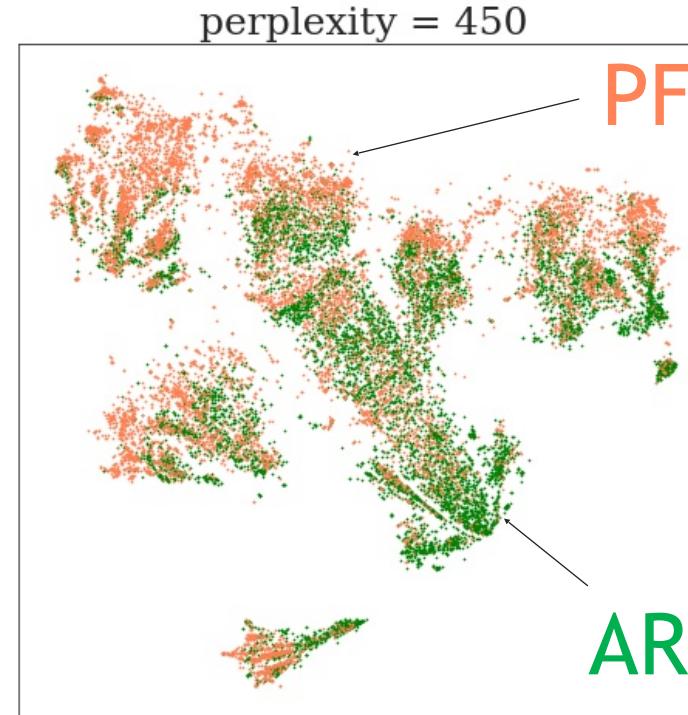
C2

- C0 (distributions)
- C1 (PCA)
- C2 (t-SNE)
- C3 (neural networks)



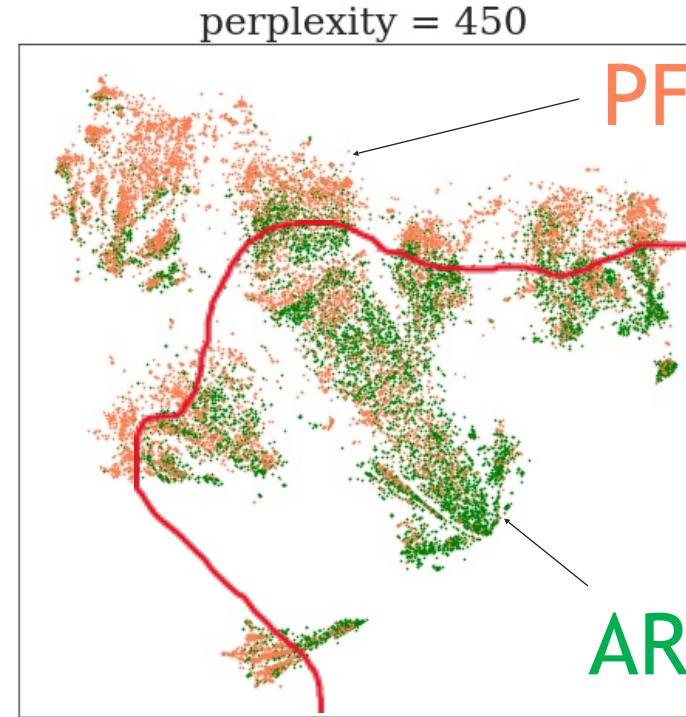
- C0 (distributions)
- C1 (PCA)
- **C2 (t-SNE)**
- C3 (neural networks)

t-SNE



- C0 (distributions)
- C1 (PCA)
- **C2 (t-SNE)**
- C3 (neural networks)

t-SNE

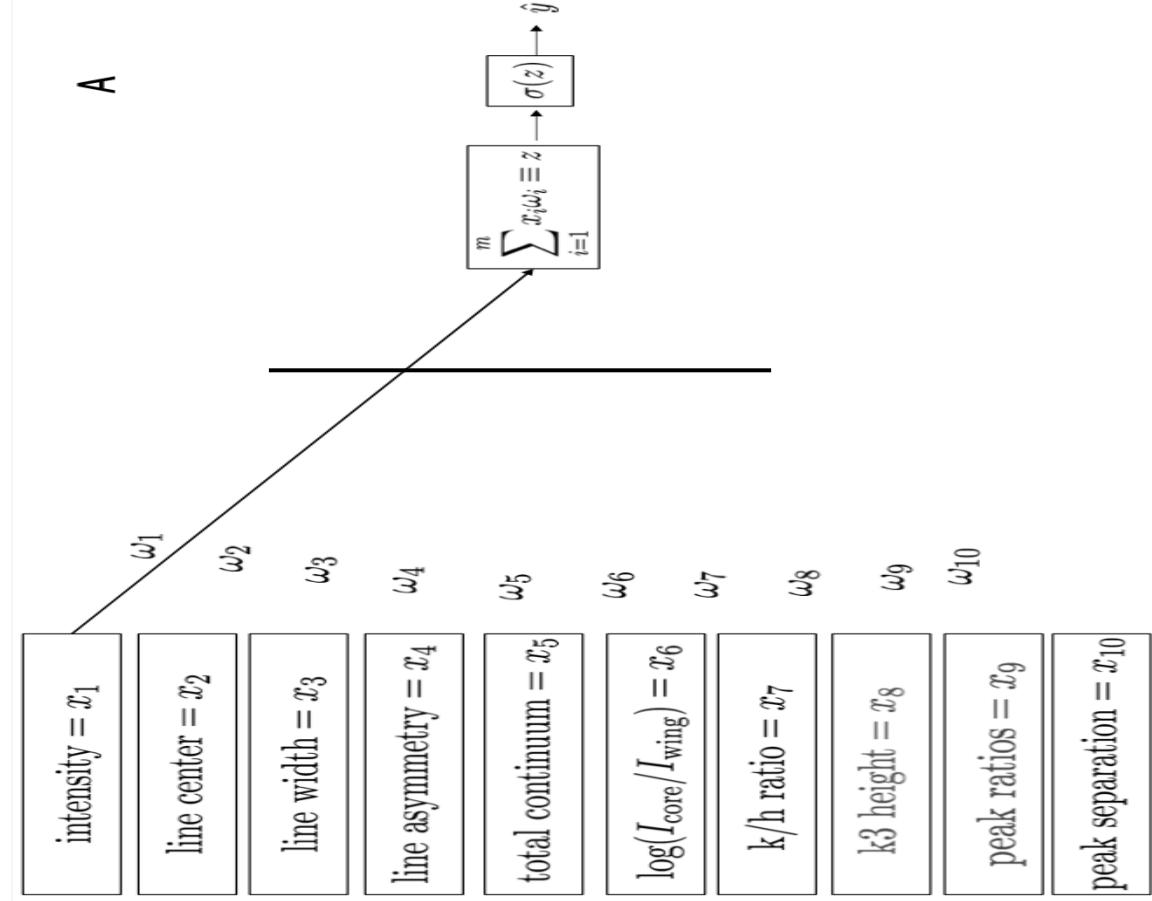


- C0 (distributions)
- C1 (PCA)
- **C2 (t-SNE)**
- C3 (neural networks)

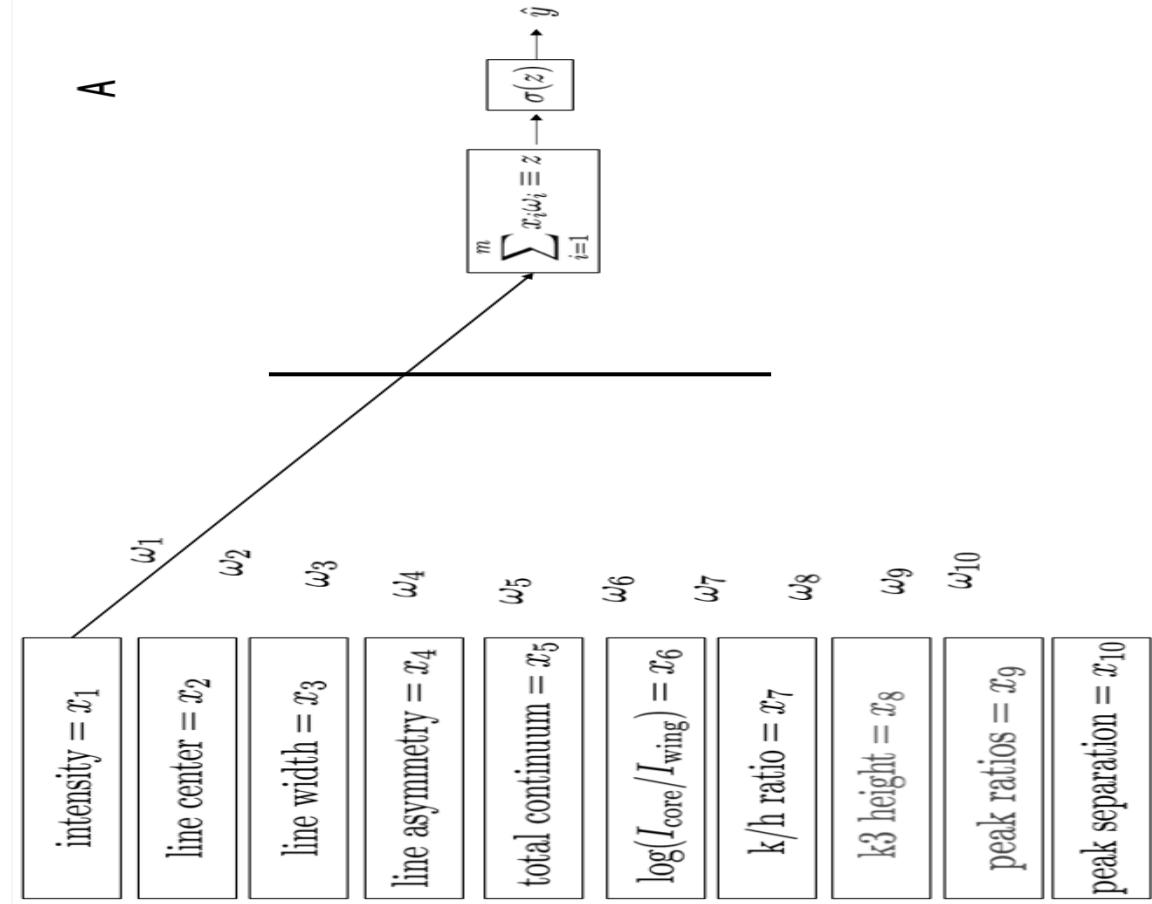
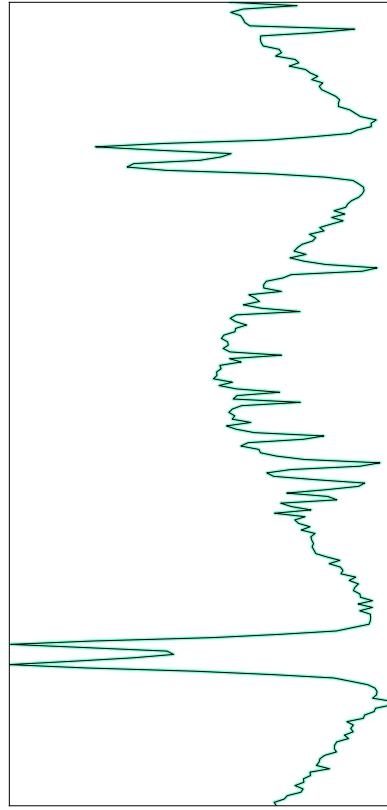
n | **w**

NN

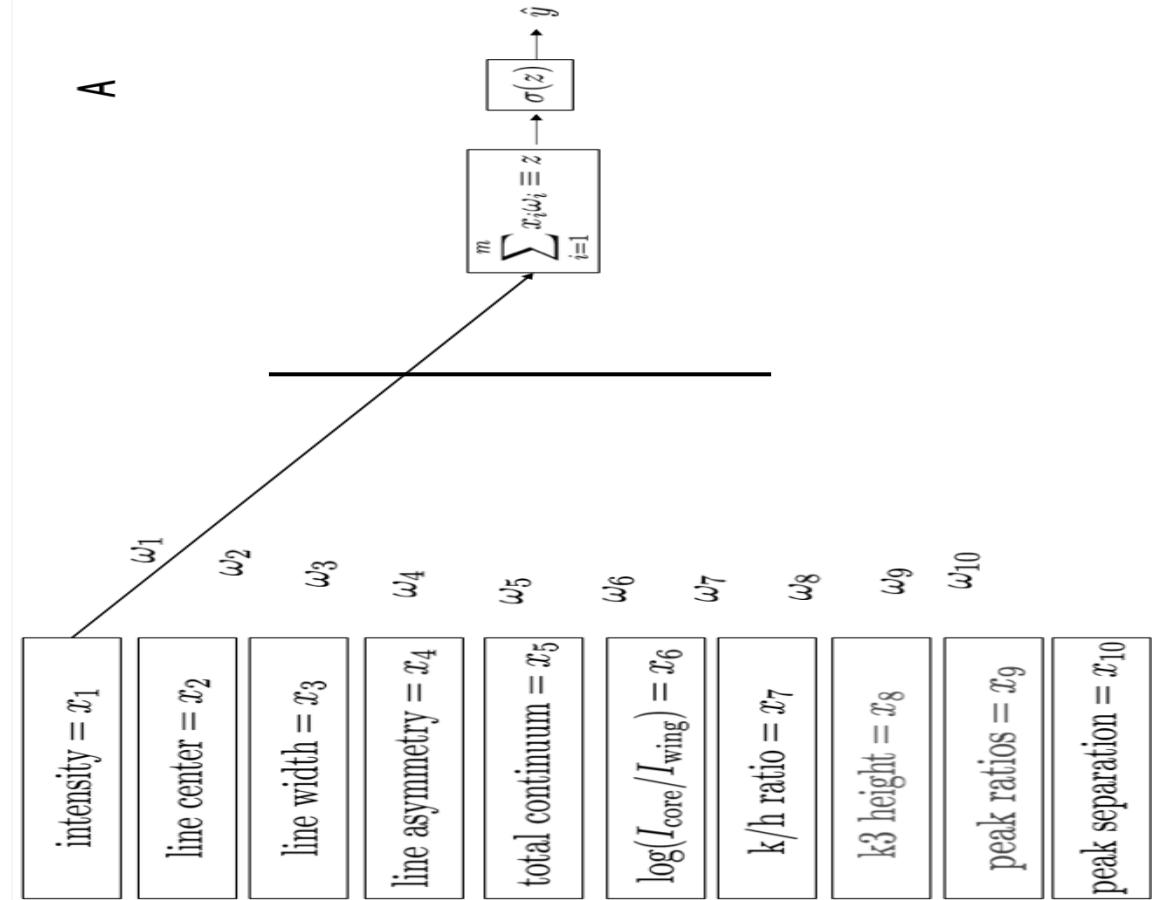
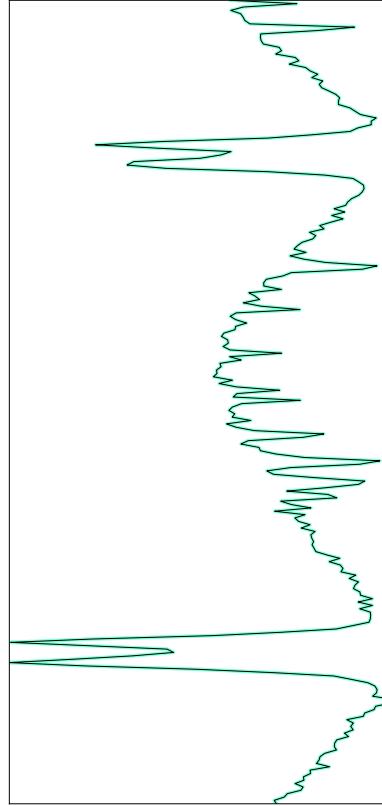
C3



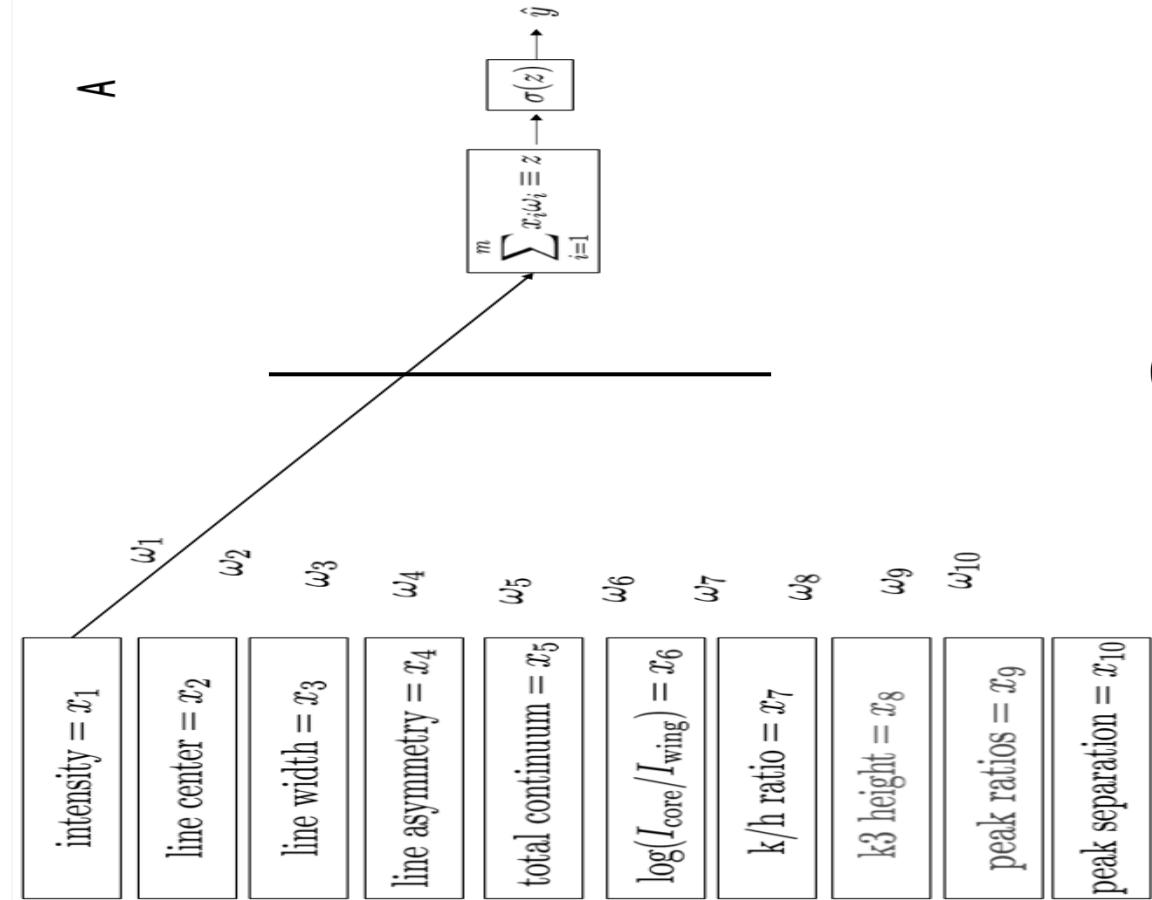
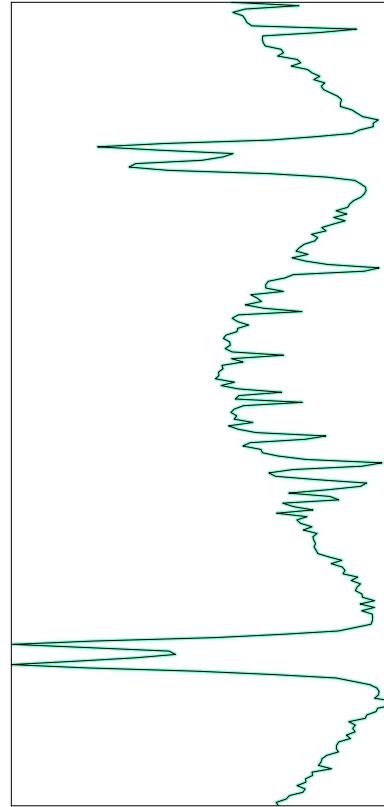
- C0 (distributions)
- C1 (PCA)
- C2 (t-SNE)
- C3 (neural networks)



- C0 (distributions)
- C1 (PCA)
- C2 (t-SNE)
- C3 (neural networks)



- C0 (distributions)
- C1 (PCA)
- C2 (t-SNE)
- C3 (neural networks)



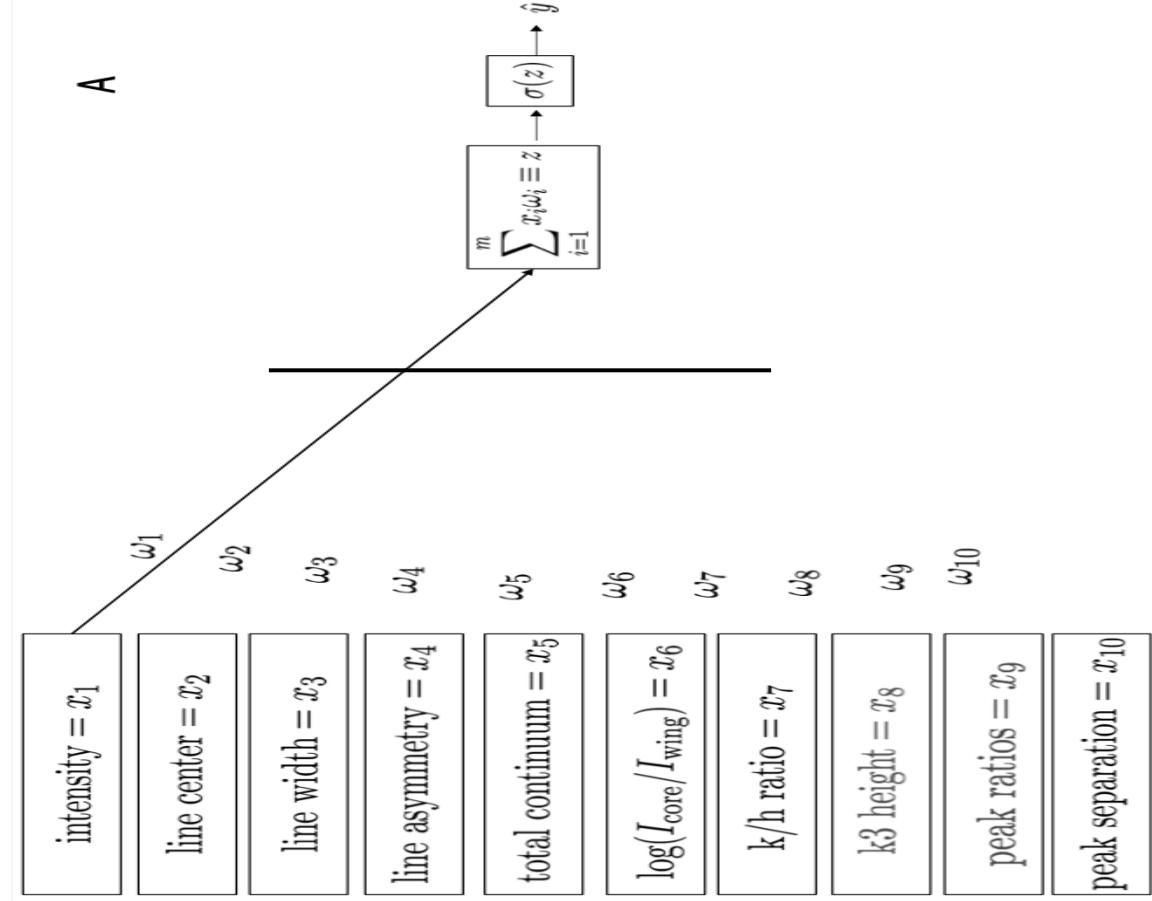
0.3

- C0 (distributions)
- C1 (PCA)
- C2 (t-SNE)
- C3 (neural networks)

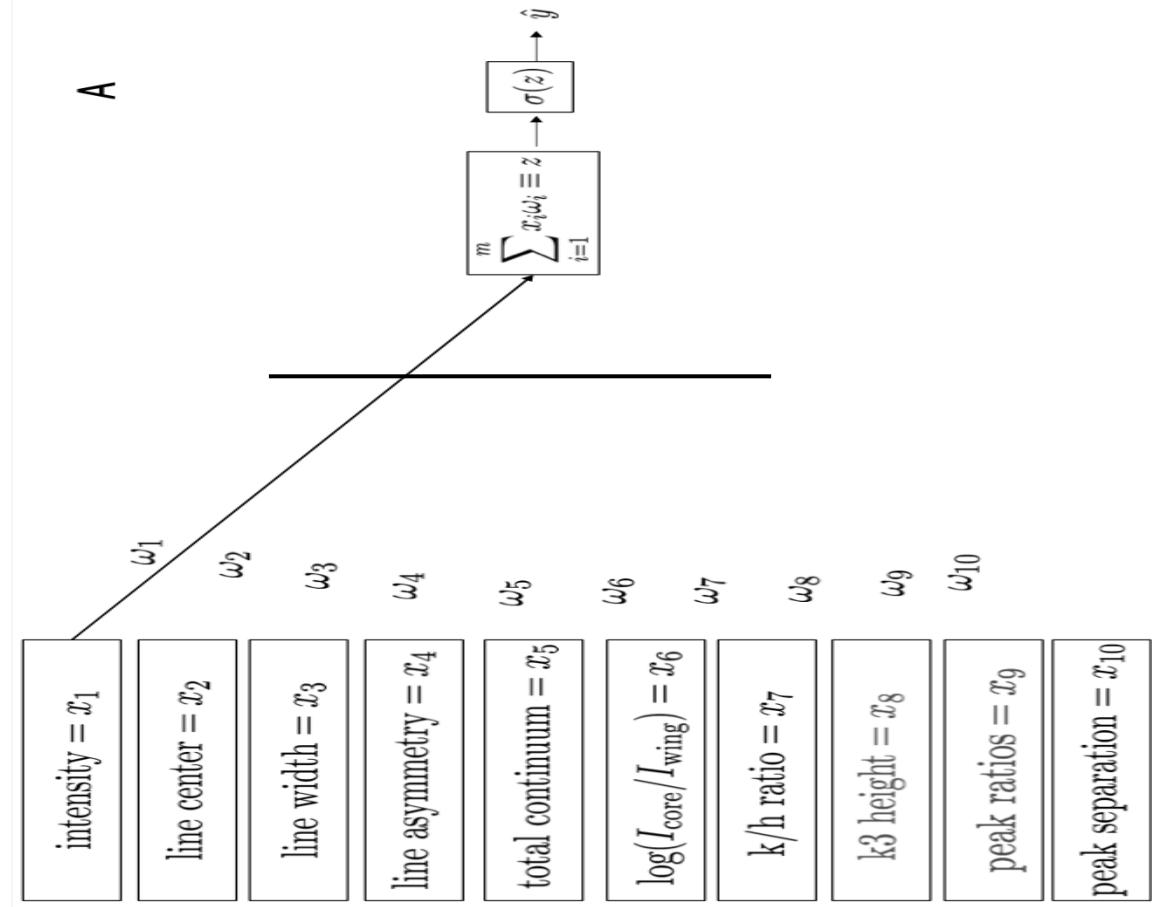
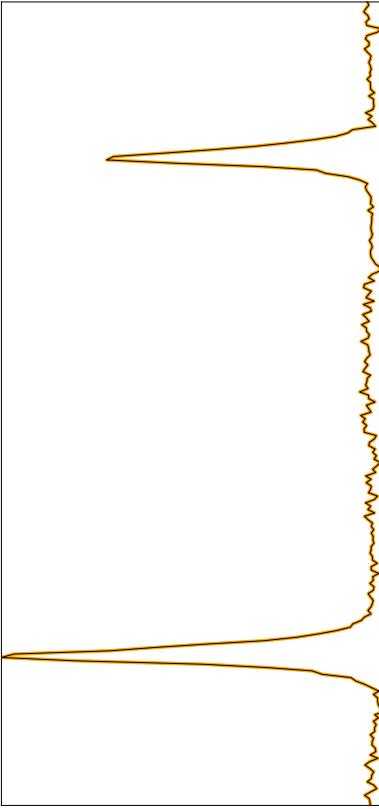
n | **w**

NN

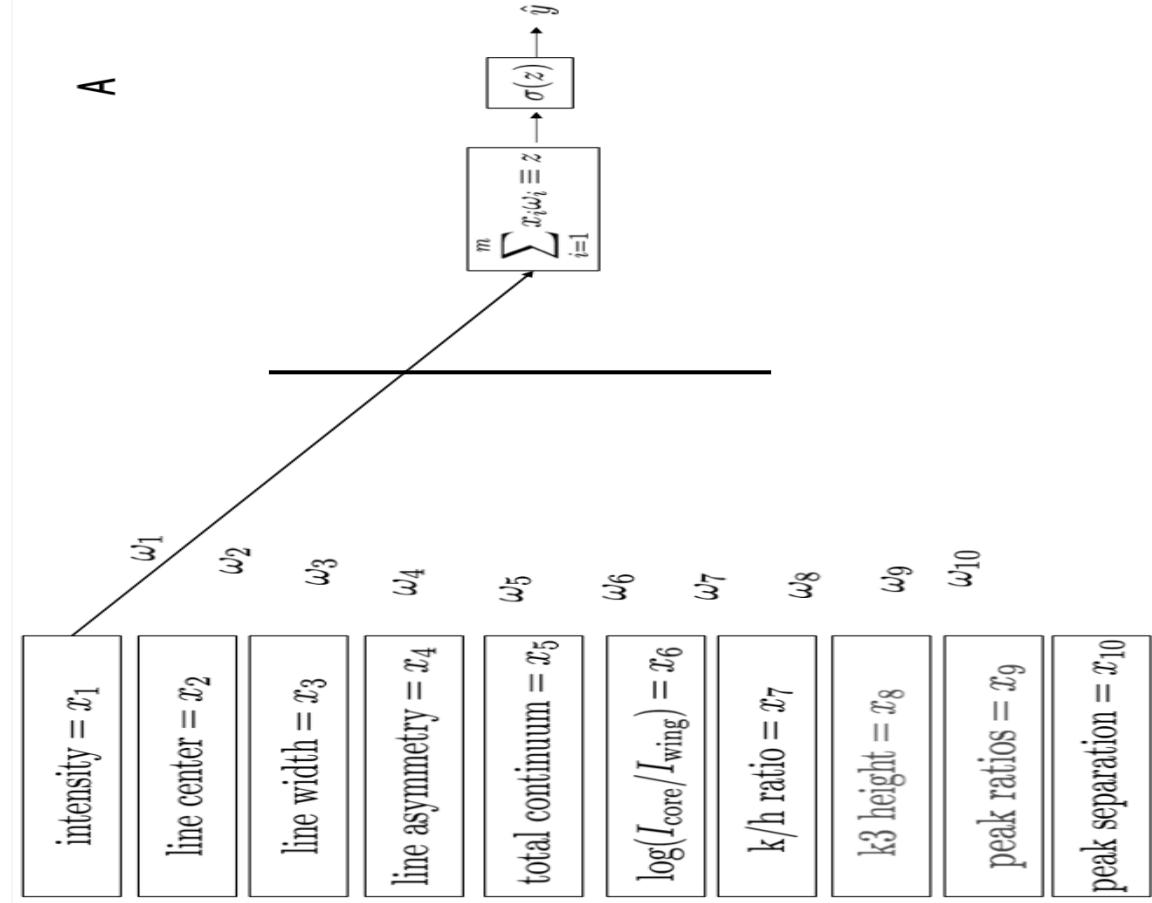
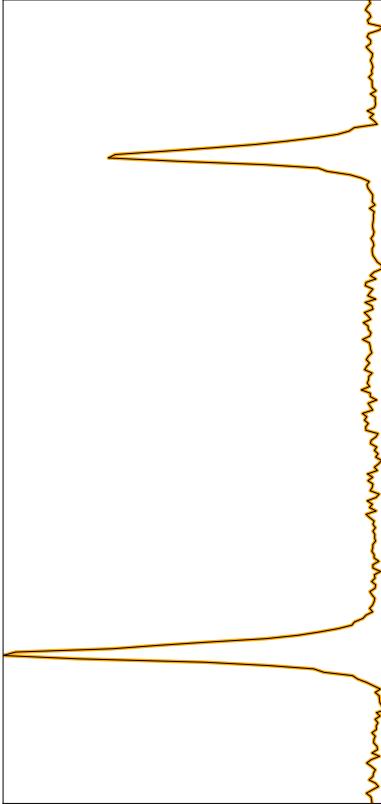
C3



- C0 (distributions)
- C1 (PCA)
- C2 (t-SNE)
- C3 (neural networks)



- C0 (distributions)
- C1 (PCA)
- C2 (t-SNE)
- C3 (neural networks)

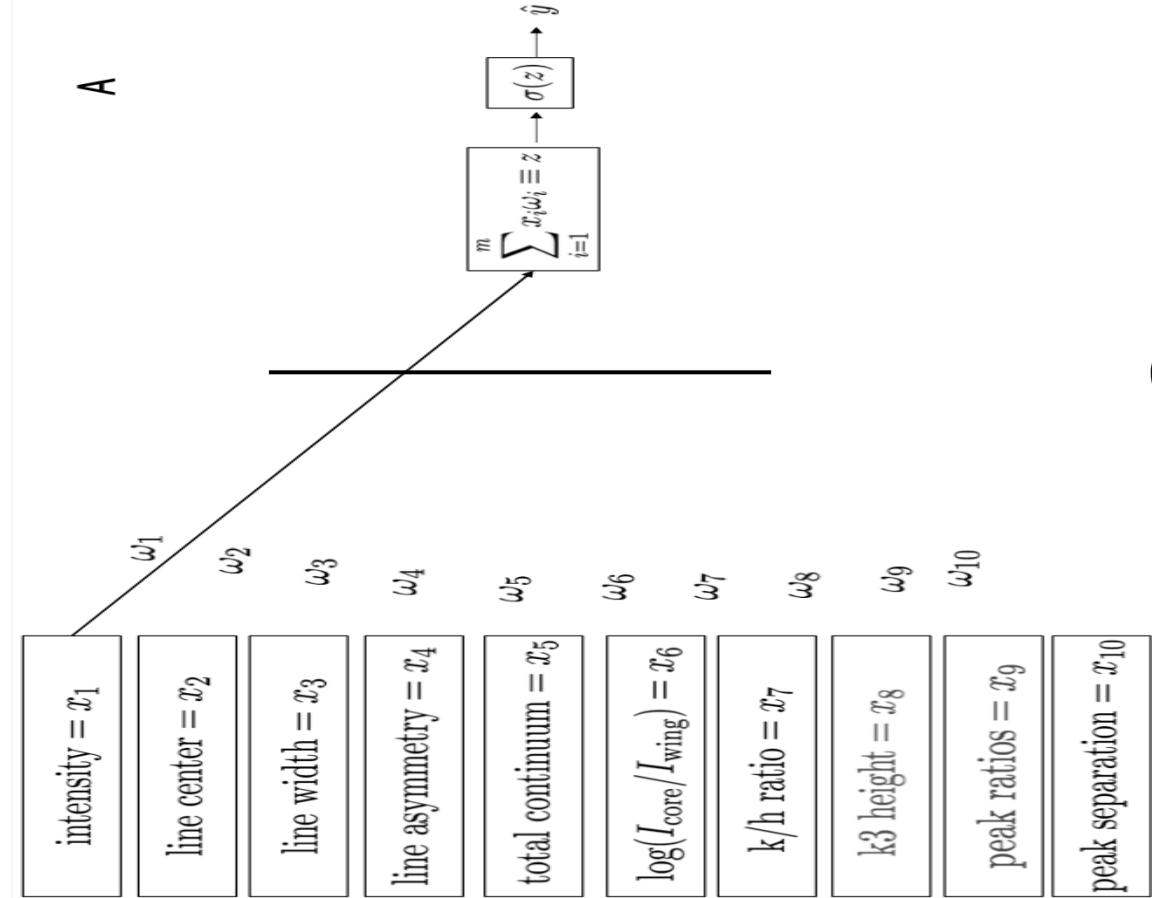
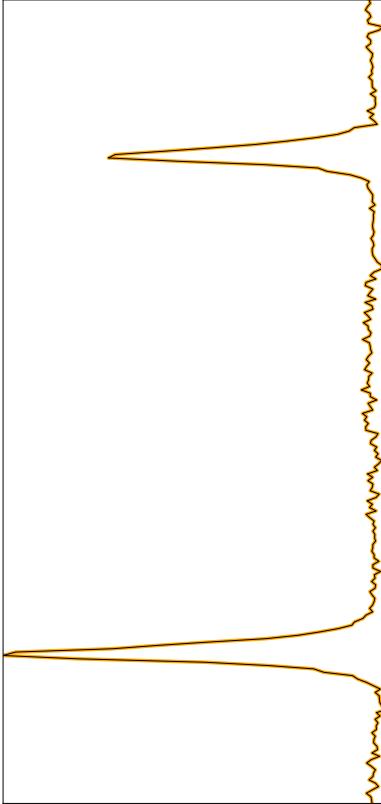


- C0 (distributions)
- C1 (PCA)
- C2 (t-SNE)
- C3 (neural networks)

n|w

NN

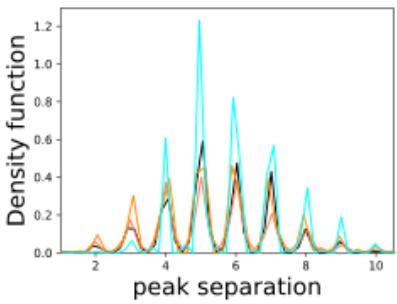
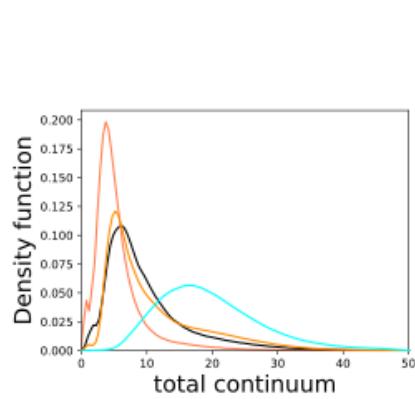
C3



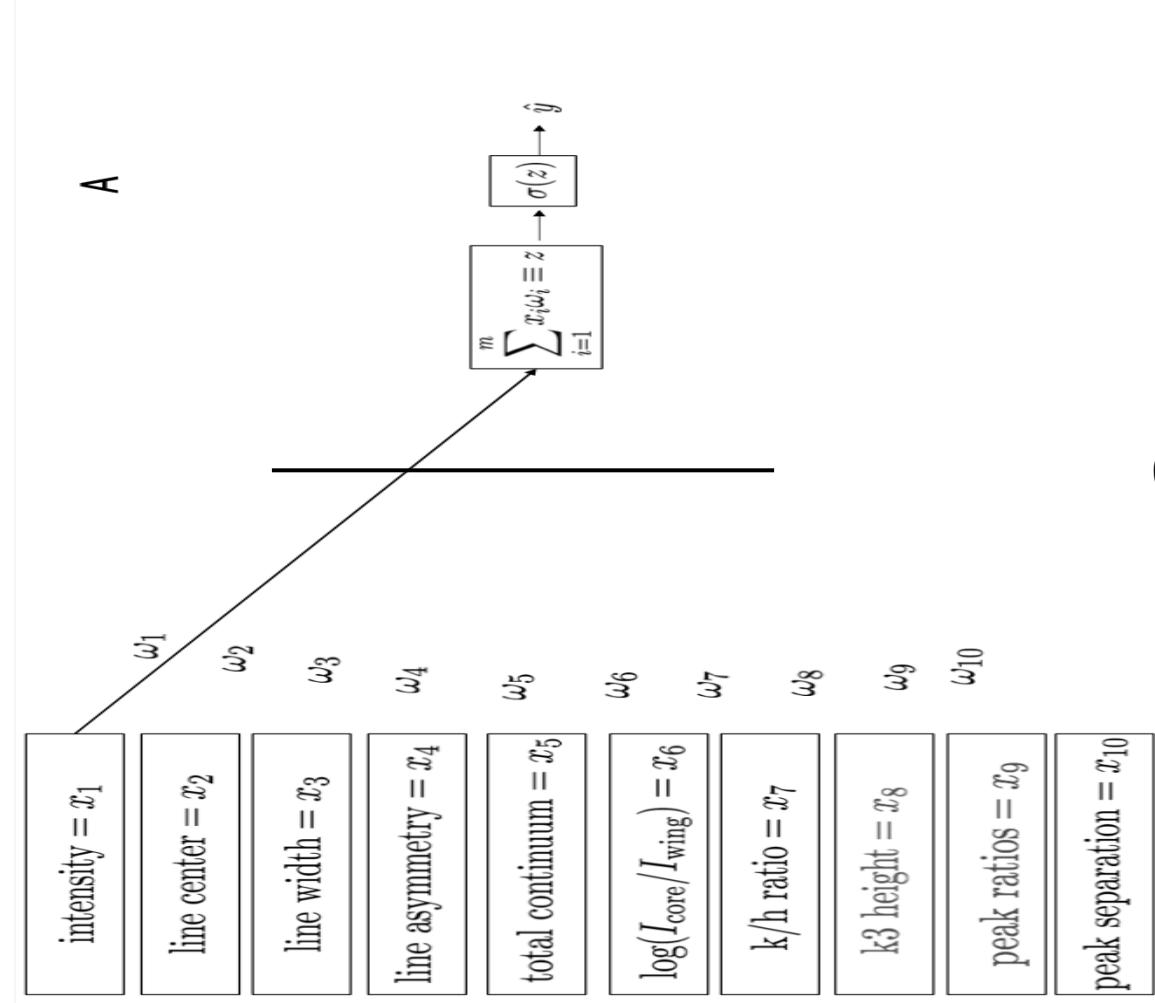
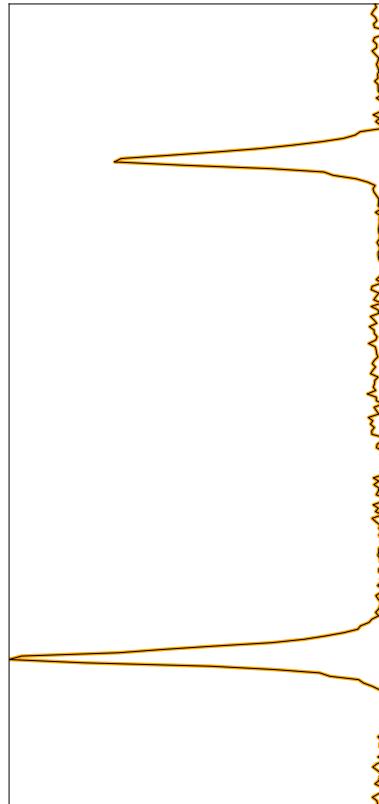
- C0 (distributions)
- C1 (PCA)
- C2 (t-SNE)
- C3 (neural networks)



C3



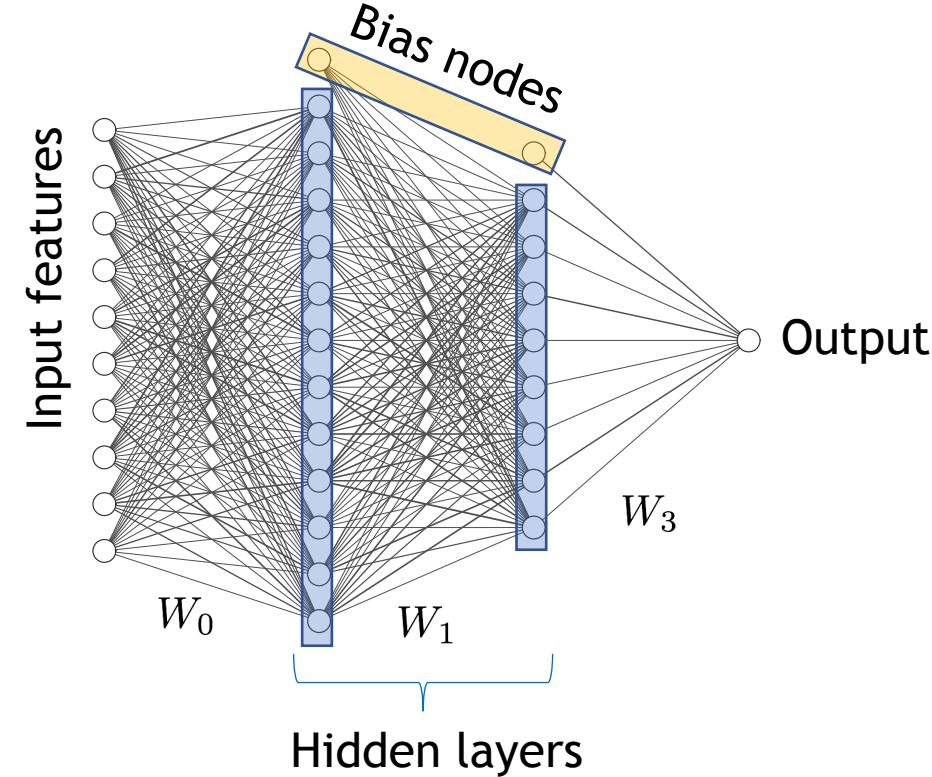
- C0 (distributions)
- C1 (PCA)
- C2 (t-SNE)
- C3 (neural networks)



0.8



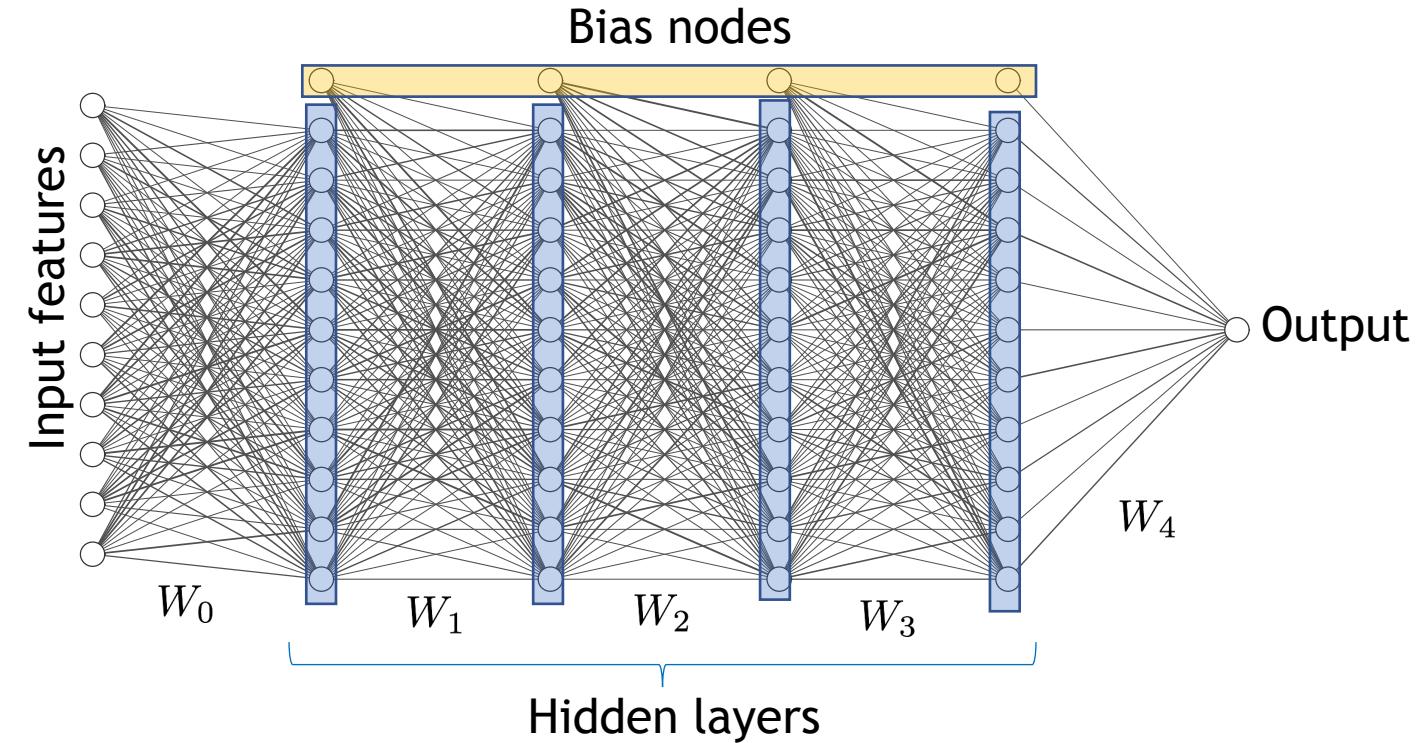
- C0 (distributions)
- C1 (PCA)
- C2 (t-SNE)
- C3 (neural networks)



- C0 (distributions)
- C1 (PCA)
- C2 (t-SNE)
- C3 (neural networks)

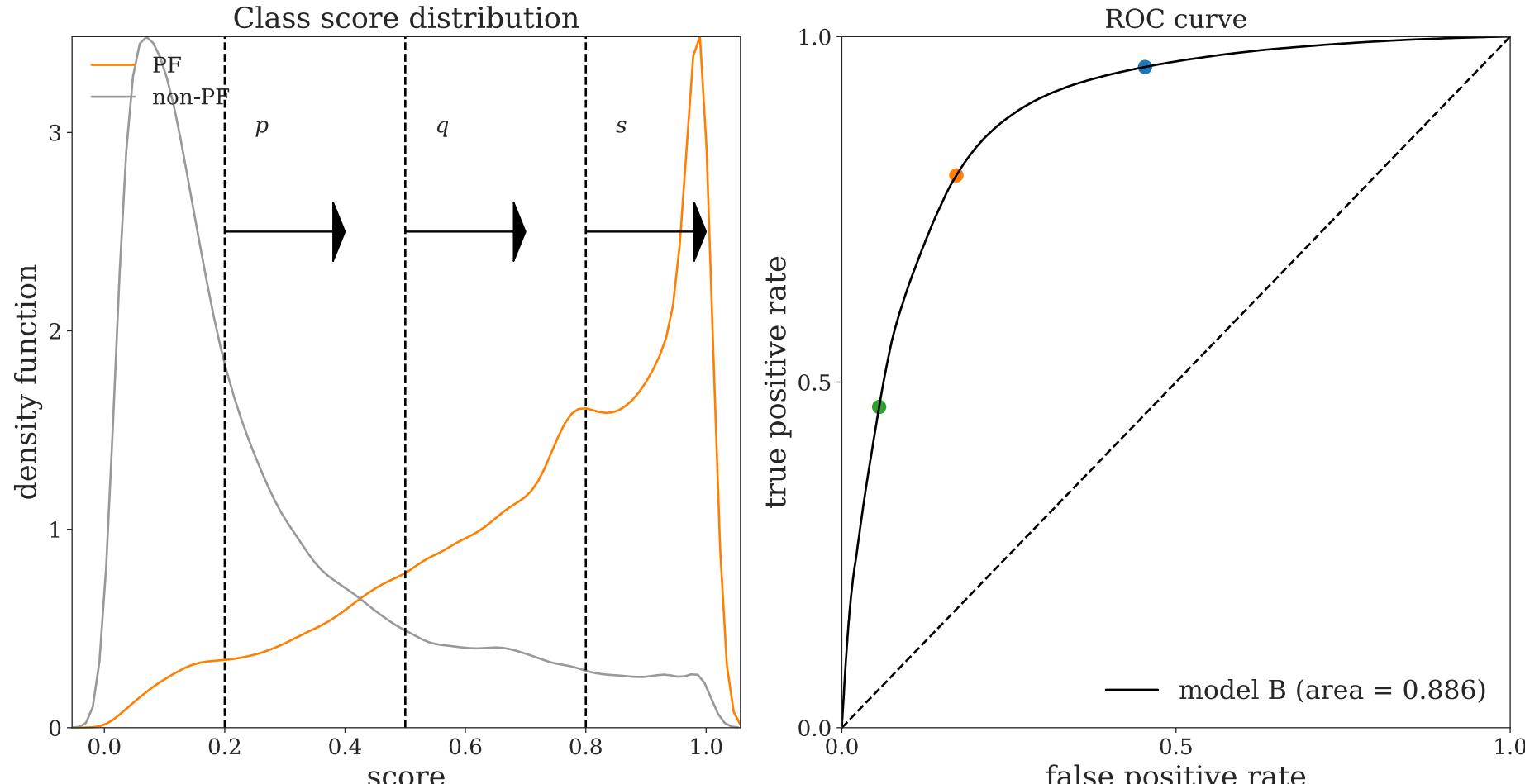


- C0 (distributions)
- C1 (PCA)
- C2 (t-SNE)
- C3 (neural networks)



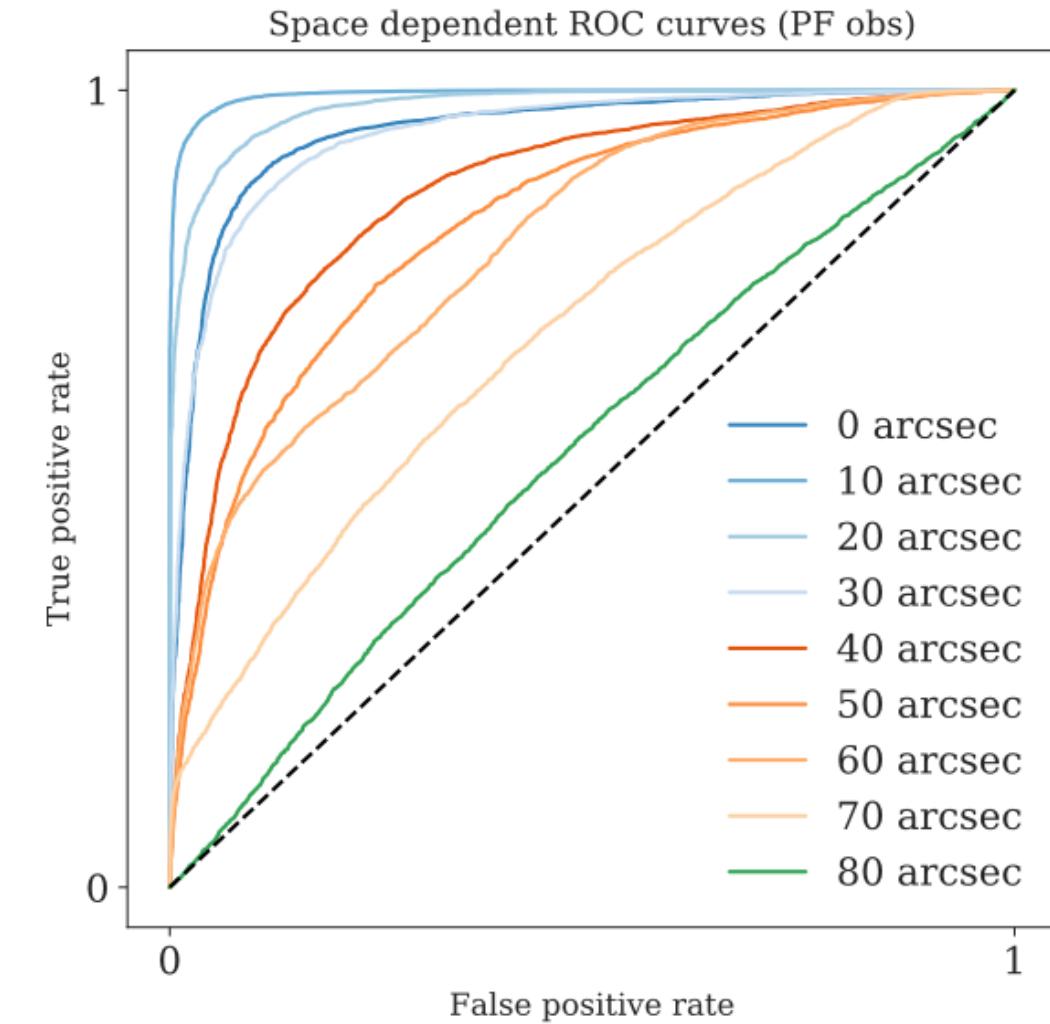
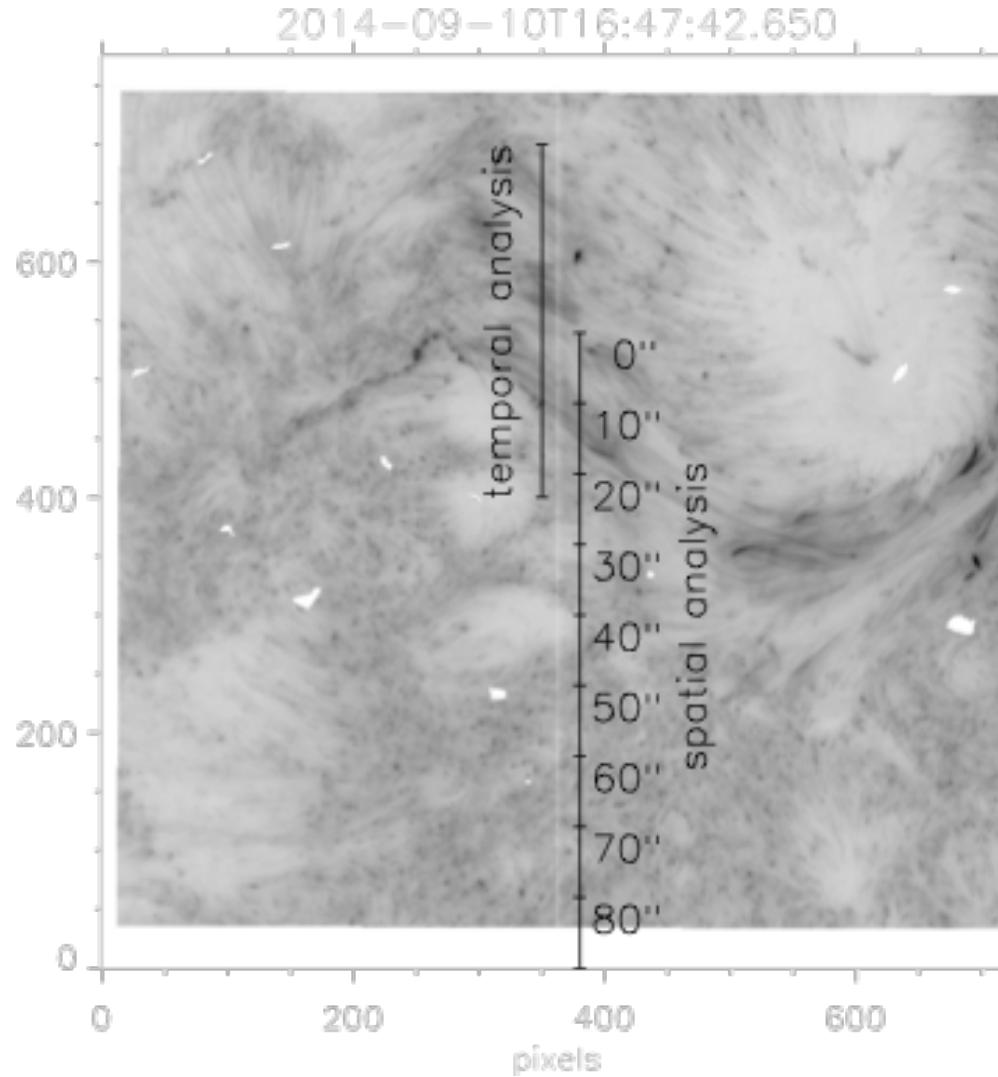
- C0 (distributions)
- C1 (PCA)
- C2 (t-SNE)
- C3 (neural networks)

Results

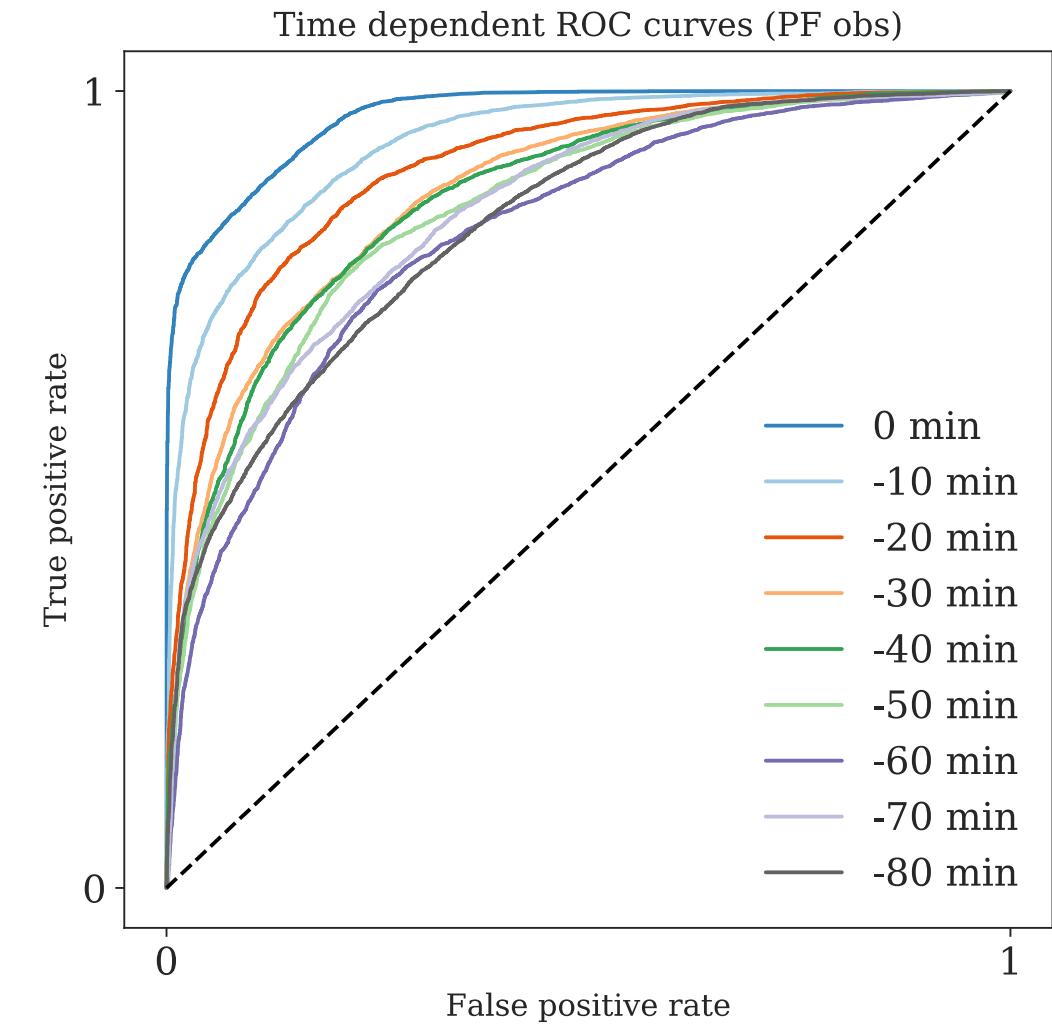
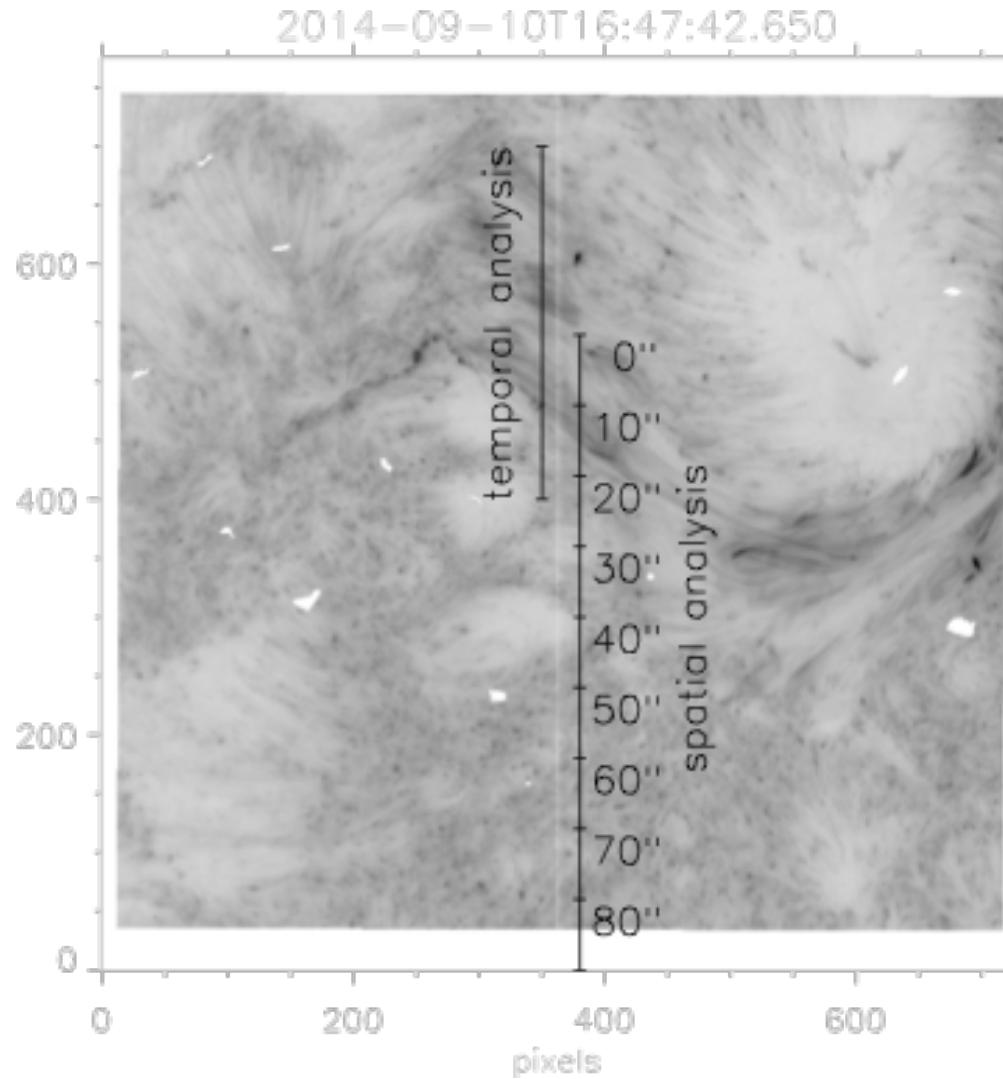


- C0 (distributions)
- C1 (PCA)
- C2 (t-SNE)
- C3 (neural networks)

Results



Results



Conclusions

- Specific solar regions produce distinct spectra
- This distinction can be leveraged by NNs to develop models that can be used as rudimentary tools for flare prediction
- Therefore, spectral data should be considered as an additional source of information alongside HMI magnetograms

Resources for solar physics and Machine learning

1. Monica Bobra and James Mason, interactive book, <http://helioml>
2. Cèdric Huwyer , IRISreader, <https://github.com/i4Ds/IRISreader>